

# A mass-loss rate determination for $\zeta$ Puppis from the quantitative analysis of X-ray emission line profiles

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## ABSTRACT

We fit every emission line in the high-resolution *Chandra* grating spectrum of  $\zeta$  Pup with an empirical line profile model that accounts for Doppler broadening and the effects of attenuation by the bulk wind. For each of sixteen lines or line complexes we find a best-fit fiducial optical depth,  $\tau_* \equiv \kappa \dot{M} / 4\pi R_* v_\infty$ , and place confidence limits on this parameter. The trend in the optical depth as a function of wavelength is completely consistent with the wavelength dependence of the atomic opacity, indicating that porosity due to large-scale clumping plays no role in explaining the only moderately asymmetric profile shapes observed in  $\zeta$  Pup. Rather, the modest signatures of attenuation in the observed line profiles are due to the low mass-loss rate, which we determine to be  $3 \pm 1 \times 10^{-6} M_\odot \text{ yr}^{-1}$  from fitting the ensemble of optical depth values derived from each line. The largest source of uncertainty in the derived mass-loss rate is the uncertainty in the elemental abundances for  $\zeta$  Pup, which affects the continuum opacity of the wind.

**Key words:** stars: early-type – stars: mass-loss – stars: winds, outflows – stars: individual:  $\zeta$  Pup – X-rays: stars

## 1 INTRODUCTION

Massive stars can lose a significant fraction of their original mass during their short lifetimes due to their strong, radiation-driven stellar winds. Accurate determinations of these stars' mass-loss rates are therefore important from an evolutionary point of view, as well as for understanding the radiative driving process itself. Massive star winds are also an important source of energy, momentum, and (chemically enriched) matter deposition into the interstellar medium, making accurate mass-loss rate determinations important from a galactic perspective.

A consensus appeared to be reached by the late 1990s that the mass-loss rates of O stars were accurately known observationally and theoretically, using the modified (Pauldrach et al. 1986) CAK (Castor et al. 1975) theory of radiation-driven stellar winds. This understanding was thought to be good enough that UV observations of spectral signatures of their winds could be used to determine

luminosities with sufficient accuracy to make extragalactic O stars standard candles (Puls et al. 1996).

This consensus has unraveled in the last few years, mostly from the observational side, where a growing appreciation of wind clumping – an effect whose importance has long been recognized (Hillier & Miller 1999) (*add refs for Hamann; Moffat work on clumping*) – has led to a re-evaluation of mass-loss rate diagnostics, including H $\alpha$  emission, radio and IR free-free emission, and UV absorption (Bouret et al. 2005; Puls et al. 2006; Fullerton et al. 2006). Accounting for small-scale clumping that affects density squared emission diagnostics – and also ionization balance and thus ionic column density diagnostics like UV resonance lines – leads to a downward revision of mass-loss rates by a factor of several, with a fair amount of controversy over the actual factor (Hamann et al. 2008).

X-ray emission line profile analysis provides a good and independent way to measure the mass-loss rates of O stars. Like the UV absorption line diagnostics, X-ray emission profile diagnostics are sensitive to the wind column density and thus are not directly affected by clumping, as the density-

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squared diagnostics are. Unlike the UV absorption line diagnostics, however, X-ray profile analysis is not very sensitive to the ionization balance and as it relies on continuum opacity rather than line opacity, it is not subject to the uncertainty associated with saturated absorption lines that hamper the interpretation of the UV diagnostics.

In this paper, we apply a quantitative line profile analysis to the *Chandra* grating spectrum of the early O supergiant,  $\zeta$  Pup, the nearest O star to the Earth and a star that has long been used as a canonical example of an early O star with a strong radiation-driven wind. Previous analysis of the same *Chandra* data has established that the kinematics of the X-ray emitting plasma, as diagnosed by the line widths, are in good agreement with wind-shock theory, and that there are modest signatures of attenuation of the X-rays by the dominant cold wind component in which the shock-heated X-ray emitting plasma is embedded (Kramer et al. 2003).

The work presented here goes beyond the profile analysis reported in that paper in several respects. We analyze many lines left out of the original study that are weak, but which carry a significant amount of information. We better account for line blends and are more careful to exclude those lines where blending cannot be adequately modeled. We model the continuum emission underlying each line separately from the line itself. We use a realistic model of the spectrometers' responses and the telescope and detector effective area. And we include the High Energy Grating (HEG) spectral data, where appropriate, to augment the higher signal-to-noise Medium Energy Grating (MEG) data that Kramer et al. (2003) reported on.

Implementing all of these improvements enables us to derive highly reliable values of the fiducial wind optical depth parameter,  $\tau_* \equiv \kappa \dot{M} / 4\pi R_* v_\infty$ , for each of sixteen emission lines or line complexes in the *Chandra* grating spectrum of  $\zeta$  Pup. Using a custom-computed model of the wavelength-dependent wind opacity,  $\kappa$ , and values for the star's radius,  $R_*$ , and wind terminal velocity,  $v_\infty$ , derived from UV and optical observations, we can fit a value of the mass-loss rate,  $\dot{M}$ , to the ensemble of  $\tau_*$  values, and thereby determine the mass-loss rate of  $\zeta$  Pup based on the observed X-ray emission line profiles.

In doing this, we also can verify that the wavelength-dependence of the profile-dependent optical depth values is consistent with that of the atomic opacity of the bulk wind, and inconsistent with a gray effective opacity. This would provide an independent line of evidence that the source of opacity in the soft X-ray is indeed the atomic opacity due to photoelectric absorption and would be in contrast to the predictions of highly porous wind models (Oskinova et al. 2006). A porosity-dominated wind would have an opacity governed by the geometric cross section of optically thick clumps. Porosity's effects can also be seen in the shapes of individual line profiles (Oskinova et al. 2006; Owocki & Cohen 2006). Our preliminary results indicate that profile models that include porosity are not favored over ones that do not (Cohen et al. 2008). We will extend this result in a forthcoming paper but do not address the effect of porosity on line profile shapes directly in the current paper. *Coauthors: should we keep these last three sentences about porous profiles? We address this briefly in the discussion section.*

The paper is organized as follows: We begin by describ-

ing the *Chandra* data set and defining a sample of well behaved emission lines for our analysis in §2. We briefly evaluate the stellar and wind properties of  $\zeta$  Pup in §3. In §4 we describe the empirical line profile model for X-ray emission lines and report on the fits to the sixteen usable lines and line complexes in the spectrum. We discuss the implications of the profile model fitting results in §5, and summarize our conclusions in §6.

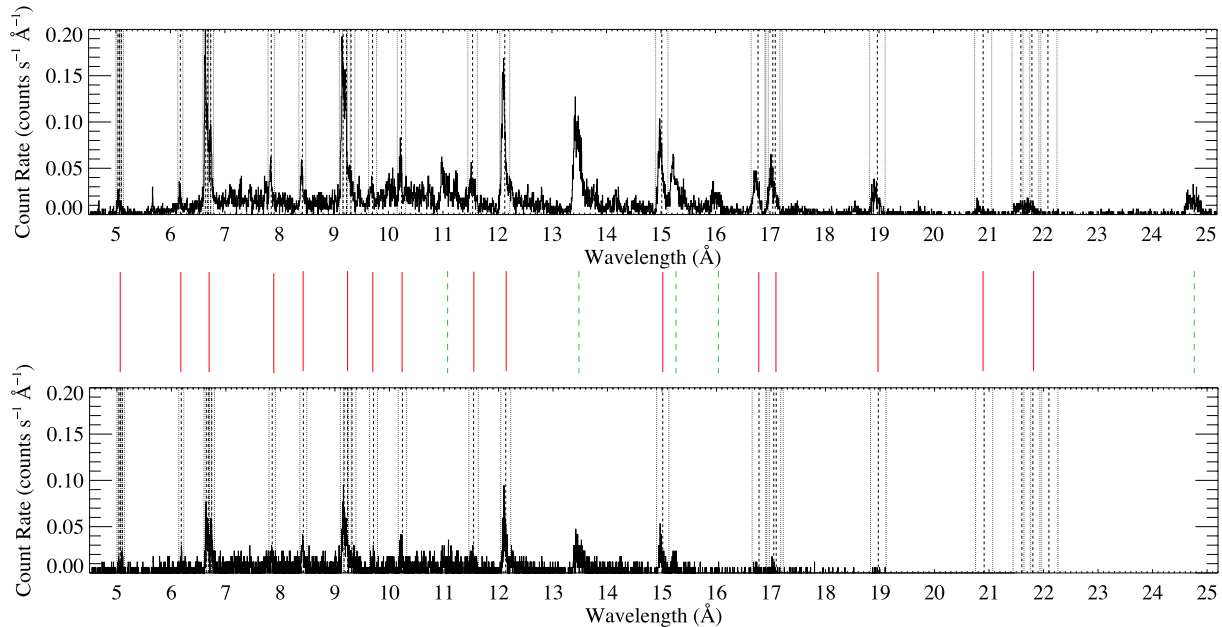
## 2 THE *Chandra* GRATING SPECTRUM

All the data we use in this paper were taken on 28-29 March 2000 in a single, 68 ks observation using the *Chandra* High-Energy Transmission Grating Spectrometer (HETGS) in conjunction with the Advanced CCD Imaging Spectrometer (ACIS) detector in spectroscopy mode. This is a photon counting instrument with an extremely low background and high spatial resolution ( $\approx 1''$ ). The first-order grating spectra we analyze have a total of 21,684 counts, the vast majority of which are in emission lines, as can be seen in Fig. 1. We modeled every line in the two spectra, as we describe in §4, and indicate in this figure which of the lines we deemed to be reliable. We only include lines in our analysis that are not so weak or severely blended that interesting parameters of the line-profile models cannot be reliably constrained.

The HETGS assembly has two grating arrays - the Medium Energy Grating (MEG) and the High Energy Grating (HEG) - with full-width half maximum (FWHM) spectral resolutions of 0.0023 Å and 0.0012 Å, respectively. This corresponds to a resolving power of  $\mathcal{R} \approx 1000$ , or a velocity of 300 km s<sup>-1</sup>, at the longer wavelength end of each grating. The wind-broadened X-ray lines of  $\zeta$  Pup are observed to have  $v_{\text{fwhm}} \approx 2000$  km s<sup>-1</sup>, and so are very well resolved by *Chandra*. The wavelength calibration of the HETGS is accurate to 50 km s<sup>-1</sup> (Marshall et al. 2004).

The two gratings, detector, and telescope assembly have significant response from roughly 2 Å to 30 Å, with typical effective areas of tens of cm<sup>2</sup>, and a strong function of wavelength. In practice, the shortest wavelength line with significant flux in the relatively soft X-ray spectra of O stars like  $\zeta$  Pup is the S xv line complex near 5 Å, and the longest wavelength line is the N VII Lyman-alpha line at 24.781 Å. The HEG response is negligible for lines with wavelengths longer than about 16 Å.

The X-ray spectrum of  $\zeta$  Pup consists of emission lines from H-like and He-like ionization stages of N, O, Ne, Mg, and Si, and numerous L-shell lines of iron, primarily Fe XVII. The Ly $\alpha$  lines and often the  $\beta$  and even  $\gamma$  lines of the Lyman series are seen for the H-like ions. There is a weak bremsstrahlung continuum beneath these lines. Overall, the spectrum is consistent with an optically thin, thermal plasma in ionization equilibrium with a range of temperatures from one to several million degrees present. It is possible that there are deviations from equilibrium, although the spectrum is not of high enough quality to show this. There is some evidence from the *XMM-Newton* RGS spectrum that a few of the emission lines are optically thick (Leutenegger et al. 2007); a possibility we will take into account when discussing the results for those lines.



**Figure 1.** The entire usable portions of the MEG (top) and HEG (bottom) first order (negative and positive orders coadded) spectra of  $\zeta$  Pup. The binning is native (2.5 mÅ for the HEG and 5 mÅ for the MEG). Vertical dashed lines in the data panels themselves represent the laboratory rest wavelengths of important lines. The lighter dotted lines on either side of these represent the Doppler shifts associated with the star’s terminal velocity. Solid (red) vertical lines between the two spectral plots indicate the lines we successfully fit with profile models and lines we attempted to fit but which were too blended to extract meaningful model parameters are indicated by dashed (green) lines. For all blended emission lines we show only one of these solid or dashed lines. *Coauthors: Should we modify this figure? I like that it shows where the lines are, how they look in context, and what the properties are of the weak, short wavelength lines are. And how complete the analysis is. I also think the  $\pm v_\infty$  dotted lines convey information about how broadened and shifted the lines are. But it’s busy...*

### 3 THE STAR

$\zeta$  Puppis is nearby ( $d = 335^{+12}_{-11}$  pc)<sup>1</sup>, single, runaway early O supergiant (O4 If) that shows the enhanced nitrogen and deficient carbon and oxygen that is indicative of CNO cycle processed material. Helium is also overabundant (*add refs*). The star’s rapid rotation may explain the photospheric abundance anomalies, though they may instead have resulted from the supernova explosion that is invoked to explain its high space velocity (Vanbeveren et al. 1998). Detailed spectral synthesis has been carried out from the UV to the IR to determine the stellar and wind properties of  $\zeta$  Pup, which we list in Table 1. Most of these are taken from Puls et al. (2006). There is a range of wind property determinations in the extensive literature on  $\zeta$  Pup. The terminal velocity of the wind may be as low as 2200 km s<sup>-1</sup> (Lamers & Leitherer 1993), and as high as 2485 km s<sup>-1</sup> (Prinja et al. 1990), though we consider the determination by the Munich group (Puls et al. 2006), of 2250 km s<sup>-1</sup>, to be the most reliable. Mass-loss rate determinations vary as well, partly because of the uncertainty in the distance, but also because each mass-loss rate diagnostic is subject to uncertainty: density squared diagnostics like H $\alpha$  and free-free

emission are affected by clumping, no matter the size scale or optical depth of the clumps. Mass-loss rates from UV absorption lines are subject to uncertain ionization corrections. In the last few years there have been attempts to account for clumping when deriving mass-loss rates from both density-squared diagnostics and UV absorption diagnostics. We list several mass-loss rate determinations in the table, with notes about the assumptions behind each determination. The X-ray line profile diagnostics of mass-loss rate that we employ in this paper are not directly affected by clumping; although very large scale porosity can affect the profiles (Oskinova et al. 2006; Owocki & Cohen 2006).

The star shows periodic variability in various UV wind lines (*ref MEGA project*) as well as H $\alpha$  (Berghoefter et al. 1996). Its broad-band X-ray properties are normal for an O star, with  $L_x \approx 10^{-7} L_{\text{Bol}}$  and a soft spectrum (Hillier et al. 1993), dominated by optically thin thermal line and free-free emission from plasma with a temperature of a few million degrees. The emission measure filling factor of the wind is small, roughly one part in 10<sup>3</sup>. Weak soft X-ray variability, with an amplitude of 6 percent, and a period consistent with the 18 hr H $\alpha$  period, was detected with *ROSAT* (Berghoefter et al. 1996). This low-level variability appears not to affect the *Chandra* data.

<sup>1</sup> The original Hipparcos distance determination (Perryman et al. 1997) had rather large error bars ( $429^{+120}_{-77}$  pc); while the value we quote in the table is from a recent reanalysis of the data (van Leeuwen 2007a,b). The derived stellar parameters and mass-loss rate depend on the distance. And the distance also has implications for the determination of the birth place of  $\zeta$  Pup.

**Table 1.** Stellar and wind parameters adopted from Puls et al. (2006)

parameter	value
Mass <sup>a</sup>	53.9 M <sub>⊙</sub>
T <sub>eff</sub>	39000 K
R <sub>*</sub>	18.6 R <sub>⊙</sub>
v <sub>rot</sub> sin i <sup>b</sup>	230 km s <sup>-1</sup>
v <sub>∞</sub>	2250 km s <sup>-1</sup>
β	0.9
$\dot{M}^c$	8.3 × 10 <sup>-6</sup> M <sub>⊙</sub> yr <sup>-1</sup>
$\dot{M}^d$	4.2 × 10 <sup>-6</sup> M <sub>⊙</sub> yr <sup>-1</sup>
$\dot{M}^e$	1.5 × 10 <sup>-6</sup> M <sub>⊙</sub> yr <sup>-1</sup>

<sup>a</sup> From Repolust et al. (2004).

<sup>b</sup> From Glebocki et al. (2000).

<sup>c</sup> Unclumped value from Puls et al. (2006).

<sup>d</sup> Also from Puls et al. (2006), but the minimum clumping model, in which the far wind, where the radio emission arises, is unclumped, but the inner wind, where the H $\alpha$  is produced is clumped. Note that the methodology of Puls et al. (2006) only enables a determination to be made of the *relative* clumping in different regions of the wind.

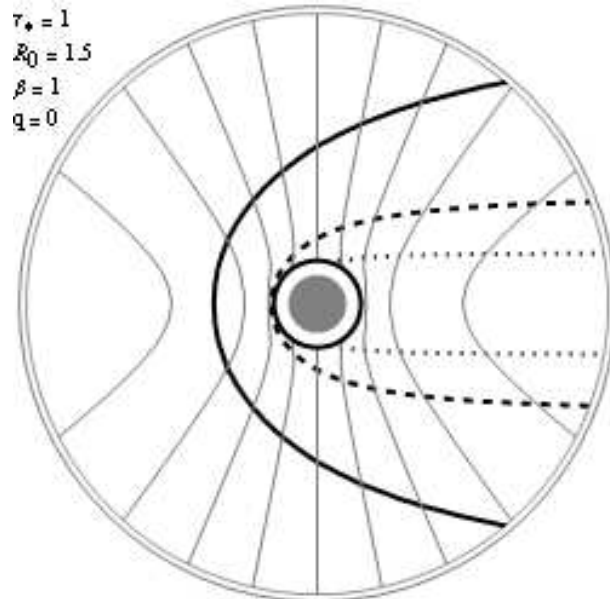
<sup>e</sup> From Bouret et al. 2008 (*Note to coauthors: This is from JC's Kauai poster; nothing seems to have appeared yet; J & J, you did this work with JC? And a paper is in the works?*), from detailed UV spectral modeling, assuming clumping.

## 4 EMISSION LINE PROFILE MODEL FITTING

### 4.1 The Model

The X-ray emission line profile model we fit to each line was first described by Owocki & Cohen (2001). It is a simple, spherically symmetric model that assumes that the local emission scales as the ambient density squared and that the many sites of hot, X-ray emitting plasma are smoothly distributed throughout the wind above some onset radius,  $R_o$ , which is expected to be several tenths of a stellar radius above the photosphere in the line-driven instability scenario (Owocki et al. 1988; Feldmeier et al. 1997; Runacres & Owocki 2002). Attenuation of the emitted X-rays occurs in the bulk, cool ( $T \approx T_{\text{eff}}$ ) wind component via photoelectric absorption, mainly out of the inner shell of elements N through Si and also out of the L-shell ( $n = 2$ ) of Fe. We assume that the atomic opacity of the cool wind, while a function of wavelength, does not vary significantly with radius. We further assume a beta-velocity law,  $v = v_\infty(1 - R_*/r)^\beta$ , for both wind components, with  $\beta = 1$  and  $v_\infty = 2250$  km s<sup>-1</sup> as given by UV observations (Puls et al. 2006). The local velocity controls the wavelength dependence of the emissivity, the local optical depth governs the wavelength-dependent attenuation, and the density affects the overall level of emission. The first two of these effects can be visualized in Fig. 2.

We cast the expression for the line profile first in spherical coordinates, but cast some of the quantities explicitly in terms of ray coordinates, with the origin at the center of the star and the observer at  $z = \infty$ . We integrate the specific intensity along rays of given impact parameter,  $p$ , and then



**Figure 2.** A visualization of the wind Doppler shift and optical depth – the two effects that govern the observed, broadened and asymmetric line shapes. The observer is on the left, and the light solid contours represent the line-of-sight velocity in increments of  $0.2v_\infty$ , with the blue shifts arising in the left hemisphere and the red shifts in the right. The star is the gray circle at the center, and the inner radius of the wind X-ray emission,  $R_o$ , is indicated at  $1.5 R_*$  by the solid black circle. The solid heavy contour represents the locus of points with optical depth  $\tau = 0.5$ , and the dashed and dotted contours represent  $\tau = 1$  and  $2$ , respectively. The model parameters visualized here are nearly identical to those of the best-fitting model for the Mg XII Ly $\alpha$  line shown in Fig. 7. *Note to coauthors: We will clean this figure up, removing the parameter labels in the upper left, incorporating them into the caption, and perhaps labeling some of the contours.*

integrate over rays. Integrating over the volume of the wind, we have:

$$L_\lambda = 8\pi^2 \int_{-1}^{+1} d\mu \int_{R_o}^{\infty} \eta_\lambda(\mu, r) r^2 e^{-\tau(\mu, r)} dr, \quad (1)$$

where  $L_\lambda$  is the luminosity per unit wavelength – it is the X-ray line profile. The angular coordinate  $\mu \equiv \cos \theta$ , and  $\eta_\lambda$  is the wavelength-dependent emissivity that accounts for the Doppler shift of the emitting parcel of wind material (which is completely determined, under the assumptions of spherical symmetry and the velocity law, according to its location,  $(\mu, r)$ ). The emissivity has an additional radial dependence due to the fact that it is proportional to the square of the ambient plasma density. The optical depth,  $\tau$ , is computed along a ray,  $z = \mu r$ , for each value of the impact parameter,  $p = \sqrt{1 - \mu^2} r$ , as

$$\tau(\mu, r) = t(p, z) = \int_z^\infty \kappa \rho(r') dz', \quad (2)$$

where the dummy radial coordinate is given by  $r' \equiv \sqrt{z'^2 + p'^2}$ . The opacity,  $\kappa$ , does not vary significantly across a line (recall it is due to continuum processes – the strong wavelength dependence across a line profile arises purely from the geometry indicated in Fig. 2). Using the continuity equation and the beta-velocity law of the wind, we have:

$$t(p, z) = \tau_* \int_z^\infty \frac{R_* dz'}{r'^2 (1 - R_*/r')^\beta}. \quad (3)$$

We account for occultation of the back of the wind by the star by setting this optical depth integral to  $\infty$  when  $p < R_*$  and  $z < \sqrt{R_*^2 - p^2}$ . The constant at the front of eq. 3,  $\tau_* \equiv \kappa \dot{M} / 4\pi R_* v_\infty$ , is the fiducial optical depth and is equivalent to the optical depth value along the central ray, integrated down to the stellar surface, in the case where  $v = v_\infty$ . This quantity,  $\tau_*$ , is the key parameter that describes the X-ray attenuation and governs the shifted and asymmetric form of the line profiles.

We note that the optical depth integral, while generally requiring numerical integration, can be done analytically for integer values of  $\beta$ . We use  $\beta = 1$  throughout this paper (though we report on tests we did for non-integer  $\beta$  values in §4.3), and for that value of the parameter, the optical depth integral evaluates as:

$$t(p > 1, z) = \frac{R_* \tau_*}{z_*} \left[ \arctan \frac{R_* \mu'}{z_*} + \arctan \frac{z'}{z_*} \right]_{z' \rightarrow z}^{z' \rightarrow \infty}, \quad (4)$$

$$t(p < 1, z) = \frac{R_* \tau_*}{2z_*} \left[ \log \left( \frac{z_* - R_* \mu'}{z_* + R_* \mu'} \right) + \log \left( \frac{z_* - z'}{z_* + z'} \right) \right]_{z' \rightarrow z}^{z' \rightarrow \infty}, \quad (5)$$

where  $z_* \equiv \sqrt{R_*^2 - p^2}$ . *Coauthors: I will clean this up a bit.*

The intrinsic line profile function we assume for the emissivity at each location is a delta function that picks out the Doppler shift line resonance,

$$\eta_\lambda \propto \delta(\lambda - \lambda_o(1 - \mu v(r)/c)). \quad (6)$$

This assumption is justified because the actual intrinsic line width is dominated by thermal broadening, which is very small compared to the Doppler shift caused by the highly supersonic wind flow.

Calculating a line profile model, then, amounts to solving equations 1 and 3 for a given set of parameters,  $R_o$ ,  $\tau_*$ , the normalization (which determines the overall level of  $\eta_\lambda$ ), and an assumed wind velocity law, described by  $\beta$  and  $v_\infty$ . This last parameter influences the emissivity term through its effect on the Doppler shift as a function of radius and spherical polar angle. And for our choice of  $\beta = 1$ , eqs. 4 and 5 replace eq. 3.

The model produces broad emission lines where the overall width (in the sense of the second moment of the profile), for an assumed wind velocity law, is governed by the parameter  $R_o$ . The closer to the star's surface  $R_o$  is, the more emission there is from low-velocity wind material, which contributes to the line profile only near line center. The value of  $\tau_*$  affects the line's blue shift and asymmetry. The higher its value, the more blue shifted and asymmetric the profile. The interplay of the two parameters can be seen in figure 2 of Owocki & Cohen (2001).

## 4.2 Fitting the data

For each line in the spectrum, we first fit the empirical line-profile model, described by equations 1 and ??, to the data. Note that this model has only three free parameters: the fiducial optical depth,  $\tau_*$ , the minimum radius of X-ray emission,  $R_o$ , and the normalization of the line.

We begin the analysis procedure for each line by fitting

the weak continuum simultaneously in two regions, one to the blue side of the line and one on the red side (but excluding the wavelength range of the line itself). We assume the continuum is flat over this restricted wavelength region. We then fit the emission line over a wavelength range that is no broader than the line itself (and sometimes even narrower, due to blends with nearby lines, which can induce us to exclude contaminated portions of the line in question). The model we fit to each line is the sum of the empirical line profile model(s) we described above and the continuum model determined from the fit to the two spectral regions near the line. Note that the inclusion of the continuum does not introduce any new free parameters.

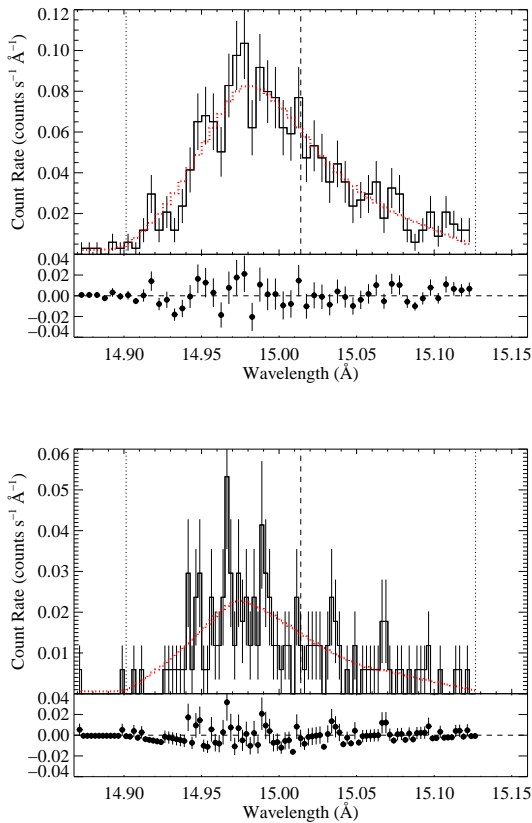
We fit the wind profile plus continuum model to both the MEG and HEG data (positive and negative first orders) simultaneously, if the HEG data are of good enough quality to warrant their inclusion (generally true only for lines with wavelengths shorter than about 16 Å), and to the MEG data only if they are not. We use the C statistic (Cash 1979) as the goodness-of-fit statistic. This is the maximum likelihood statistic for data with Poisson distributed errors, which these photon-counting X-ray spectra are. Note that the maximum likelihood statistic for Gaussian distributed data is the well-known  $\chi^2$  statistic, but it is not valid for these data, which have many bins with only a few counts, especially in the diagnostically powerful wings of the profiles.

We determine the best-fit model by minimization of the C statistic using the *fit* task in XSPEC. Once it is found, the uncertainties on each model parameter are assessed using the  $\Delta\chi^2$  formalism<sup>2</sup> outlined in chapter 15 of Press et al. (2007), which is also valid for  $\Delta C$ . We test each parameter one at a time, stepping through a grid of values and, at each step, refitting the data while letting the other model parameters be free to vary. The 68 percent confidence limits determined in this manner are what we report as the formal uncertainties in the table of fitting results, below. We also examine the confidence regions in two-dimensional subspaces of the whole parameter space in order to look for correlations among the interesting parameters.

We use the relatively strong and unblended Fe xvii line at 15.014 Å to demonstrate this fitting process. We show the MEG and HEG data for this line, along with the best-fit model (the set of model parameters,  $\tau_*$ ,  $R_o$ , and normalization that minimizes the C statistic) in Fig. 3. The best-fit model parameters are:  $\tau_* = 1.97$ ,  $R_o = 1.53 R_*$ , and a normalization of  $5.24 \times 10^{-4}$  photons  $s^{-1} \text{ cm}^{-2}$ . Using the  $\Delta C$  criterion and testing each of these parameters one at a time, we find that the 68 percent confidence limits on the fit parameters are  $1.63 < \tau_* < 2.35$ ,  $1.38 < R_o/R_* < 1.65$ , and  $5.04 \times 10^{-4} < \text{norm} < 5.51 \times 10^{-4}$ . The confidence limits should be thought of as probabilistic statements about the chance that the true parameter values lies within the given range, given the physical assumptions of the model.

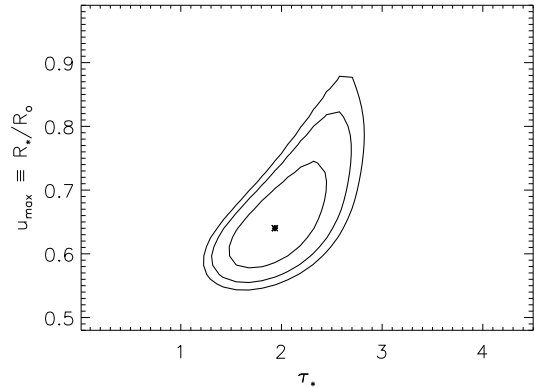
In Fig. 4 we show 68, 90, and 95 percent confidence limits in two-dimensional  $\tau_*$ ,  $R_o$  parameter space. *Note to coauthors: All the uses of the parameter  $u_{\max}$  or  $u_o$ , especially in figures such as this one, will be removed. We will discuss*

<sup>2</sup> This criterion is a specific numerical value of  $\Delta C \equiv C_i - C_{\min}$  for model realization  $i$ , where  $C_{\min}$  is the C statistic value for the best-fit model.



**Figure 3.** The Fe XVII line at 15.014 Å in the MEG (top) and HEG (bottom), with the best-fit model superimposed. We have not done any rebinning of the data. The error bars represent Poisson, root-N, statistics. The dashed vertical line indicates the laboratory rest wavelength of the emission line, and the two dotted vertical lines indicate the wavelengths associated with the Doppler shift due to the stellar wind terminal velocity of 2250 km s<sup>-1</sup>. The model is shown as a (red) dotted histogram, while the data are shown as a (black) solid histogram with error bars. The fit residuals are shown in the horizontal windows below the data.

everything in terms of  $R_o$ , and also remake all the 2-D confidence plots in  $\tau_*, u_o$  space so they are displayed in  $\tau_*, R_o$  space. This change will happen soon. We calculate a grid of models (typically 36 by 36), optimizing the other free parameters (just the normalization, in this case) at each point in the grid, and use values of  $\Delta C = 2.30, 4.61, 6.17$  (Press et al. 2007) to define the extent of the confidence limits. Plots such as this one are a good means of examining correlations between model parameters, in terms of their abilities to produce similar features in the line profiles. We can see what the trade offs are between parameters in a quantitative way. For example, there is a slight correlation between  $u_o$  and  $\tau_*$  evident in the figure. High values of  $u_o$  ( $R_o$  close to  $R_*$ ), reduce emission on the line wing relative to the core (more emitting material at low velocity). So although high values of  $u_o$  (hot plasma as close as 1.2  $R_*$ ) are allowed at the 95 percent confidence limit, they require a large wind optical depth,  $\tau_* \approx 3$ , to compensate. High  $\tau_*$  makes lines more blue-shifted and asymmetric, increasing the emission on the line wing relative to the core. We note that the confidence

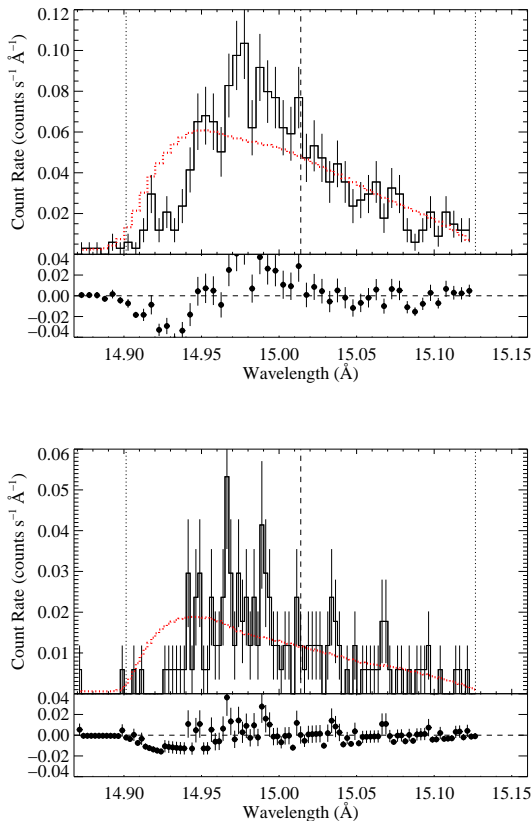


**Figure 4.** Confidence contours (68, 90, and 95 percent) for the non-porous model fitting of the the Fe XVII line at 15.014 Å. The best-fit, shown in Fig. 3, is represented by the asterisk. Note that we plot this, and all other confidence plots that involve the inner radius,  $R_o$ , in terms of  $u_{\max}$ , or  $u_o \equiv R_*/R_o$  (for now).

limits listed in the table of model fitting results, which are for individual parameters considered one at a time, will tend to differ somewhat from those inferred from these plots of joint confidence limits.

The value of  $\tau_*$  expected from the traditional mass-loss rate and a model of the wind opacity at 15 Å is  $\tau_* \approx 8$  (*Note to coauthors: The CMFGEN opacity model we show in the discussion section implies a  $\tau_*$  value closer to 7, but before redoing a bunch of fits, I'd like to make sure we're settled on the final version of our wind opacity model.*). The best-fit model with fixed  $\tau_* = 8$  is shown in Fig. 5. This model does not provide a good fit, having  $\Delta C = 108$ , implying rejection probabilities well above 99.99 percent. This is the quantitative basis for claims that the X-ray emission lines of O stars in general, and  $\zeta$  Pup in particular, are too symmetric and unshifted to be explained by the standard wind-shock scenario (Kahn et al. 2001; Cassinelli et al. 2001; Kramer et al. 2003). However, the primary goal of this paper is to quantify the mass-loss rate reduction compared to the older values from the literature by modeling the wind opacity and the effects of wind attenuation on all the line profiles simultaneously. To enable us to do this, we repeat the fitting procedure described here for the line at 15.014 Å for all of the lines and line complexes in the spectrum.

For the helium-like complexes – O VII, Ne IX, Mg XI, Si XIII, and S XV – we fit a modified version of the wind profile model in XSPEC that simultaneously fits three separate profiles with the basic parameters ( $\tau_*$  and  $R_o$ ) tied together and accounts for the altered forbidden-to-intercombination line strength ratios due to the effects of photoexcitation out of the  $2^3S$  state, which is the upper level of the forbidden line. This model, which was first described in Leutenegger et al. (2006), assumes a spatial distribution of X-ray emitting plasma, just as the basic wind profile model does, but alters the radius-dependent line ratio according to the ultraviolet mean intensity computed from an assumed model atmosphere. This model thus self-consistently accounts for the effects of the radial dependence of the individual line emissivities on both the line ratio and the profile shapes. Although the components of these complexes are blended,



**Figure 5.** The Fe xvii line at 15.014 Å in the MEG (top) and HEG (bottom), with the best-fit model having  $\tau_* = 8$  superimposed. This model fit is statistically unacceptable.

we can extract useful model parameters and confidence limits on those parameters by fitting each complex as a single entity.

We handle other line blends similarly, by fitting two (or more) separate line profile models (with an underlying continuum, fit separately) with parameters forced to be the same for each component. In some cases, like the Fe xvii lines at 17.051 and 17.096 Å, where the relative intensities of the components are well constrained by atomic physics, we obtain reliable results. In other cases, like the N vii Ly $\alpha$  line at 24.781, which is blended with the N vi He $\beta$  line, it is impossible to accurately model the relative contributions of the two lines, and no reliable information can be obtained from fitting the line blend. In still other cases, the blending is mild – through a combination of the second line being weak and the overlap region being small – and we can fit the stronger of the components reliably by simply excluding some of the data. This was the case for the Ne x Ly $\alpha$  line at 12.134 Å, where the extreme red wing is mildly blended with a weak iron line. The line blends that could not be reliably fit are indicated in Fig. 1 by the dashed vertical lines between the panels, and include the helium-like neon complex, which is blended with several iron lines.

After eliminating the complexes too blended to be reliable, we are left with sixteen lines and line complexes that could be reliably fit with the wind profile model as described in the previous subsection and as demonstrated on the Fe

xvii line at 15.014 Å, above. The results of these fits are summarized in Table 2. And we show four more representative line fits – spanning a wide range of wavelengths and derived values of  $\tau_*$  – in Figs. 6, 7, 8, and 9. We summarize the sixteen derived  $\tau_*$  and  $R_o$  values, along with their confidence limits, in Fig. 10.

### 4.3 Sensitivity of fitting results to modeling assumptions

We have made various assumptions and choices in carrying out the line-profile modeling described in the previous subsection. And we therefore have investigated many of these, again using the Fe xvii line at 15.014 Å as a test case. In this subsection, we report on the sensitivity of our results to the following assumptions and choices: background subtraction; determination of the continuum level; exclusion of portions of the line due to possible blending; inclusion of the weak HEG data; the adopted values of  $\beta$  and  $v_\infty$  for the wind; and whether to allow the X-ray volume filling factor to vary with radius (as parameterized by  $q$  in  $f \propto r^{-q}$  – see Owocki & Cohen (2001)). We will very briefly describe those factors that we found to be unimportant, and discuss in more detail those that did make a difference. The baseline model fitting we describe here is the modeling described in the previous subsection for the 15.014 Å line, except that we fit only the MEG data (so that we may evaluate the effect of including the HEG data).

We examined the default background spectra, which were very weak, and also experimented with fitting the 15.014 Å line with and without the background subtracted and found almost no difference. We therefore opt to neglect the background when fitting each of the lines in the spectrum. The sensitivity to the continuum fit is a little greater, but still nearly negligible. When we changed the continuum level by a factor of two, none of the parameter values changed by more than ten percent. Some lines in the spectrum are blended with weaker lines. The cleanest way to handle this situation is to exclude the contaminated bins from the modeling. To test the effects of this, we eliminated 0.03 Å from the red wing of the 15.014 Å line and refit the data. We then repeated this experiment eliminating 0.07 Å – leaving only about two-thirds of the data. Even in this second, extreme case, the fit parameters varied by less than ten percent and the confidence regions only expanded slightly.

For most lines, the HEG data is significantly weaker than the MEG data. We find for the 15.014 Å line that including the HEG data changes the best-fit model parameters by, at most, a few percent, but it does tighten the confidence limits somewhat. The effect of including the HEG data is more significant for the shorter wavelength lines, where the effective area of the HEG is larger. There is very little penalty for including the HEG data, so we do so for all lines shortward of 16 Å. We also fit the MEG and HEG data separately for the 15.014 line to verify that there are not systematic effects between these two spectra; and there are not. The separate fits give results that are very similar to each other, with significantly overlapping 68 percent confidence limits for all parameters.

The original Owocki & Cohen (2001) line profile model allows for a radially varying filling factor of X-ray emitting plasma, parameterized as a power law function of radius.

**Table 2.** Wind profile model fit results

ion	wavelength <sup>a</sup> (Å)	$\tau_*$	$R_o$ ( $R_*$ )	normalization <sup>b</sup> ( $10^{-5}$ ph cm <sup>-2</sup> s <sup>-1</sup> )
S xv	5.0387, 5.0648, 5.1015	0.01 <sup>+0.36</sup> <sub>-0.01</sub>	1.41 <sup>+0.15</sup> <sub>-0.11</sub>	2.56 <sup>+0.24</sup> <sub>-0.36</sub>
Si xiv	6.1822	0.49 <sup>+0.61</sup> <sub>-0.35</sub>	1.46 <sup>+0.20</sup> <sub>-0.14</sub>	0.77 <sup>+0.11</sup> <sub>-0.14</sub>
Si xiii	6.6479, 6.6866, 6.7403	0.37 <sup>+0.13</sup> <sub>-0.25</sub>	1.48 <sup>+0.05</sup> <sub>-0.04</sub>	10.9 <sup>+0.5</sup> <sub>-0.2</sub>
Mg xi	7.8503	0.65 <sup>+0.19</sup> <sub>-0.32</sub>	1.33 <sup>+0.12</sup> <sub>-0.13</sub>	1.33 <sup>+0.17</sup> <sub>-0.13</sub>
Mg xii	8.4210	1.22 <sup>+0.53</sup> <sub>-0.45</sub>	1.34 <sup>+0.18</sup> <sub>-0.21</sub>	2.95 <sup>+0.24</sup> <sub>-0.24</sub>
Mg xi	9.1687, 9.2297, 9.3143	0.92 <sup>+0.19</sup> <sub>-0.16</sub>	1.55 <sup>+0.06</sup> <sub>-0.06</sub>	17.8 <sup>+0.8</sup> <sub>-0.5</sub>
Ne x	9.7082	0.62 <sup>+1.05</sup> <sub>-0.52</sub>	1.48 <sup>+0.27</sup> <sub>-0.19</sub>	0.95 <sup>+0.15</sup> <sub>-0.15</sub>
Ne x	10.2388	1.95 <sup>+0.28</sup> <sub>-0.87</sub>	1.01 <sup>+0.00</sup> <sub>-0.00</sub>	2.99 <sup>+0.31</sup> <sub>-0.29</sub>
Ne ix	11.5440	0.83 <sup>+0.65</sup> <sub>-0.44</sub>	2.08 <sup>+0.54</sup> <sub>-0.36</sub>	5.00 <sup>+0.40</sup> <sub>-0.50</sub>
Ne x	12.1339	2.03 <sup>+0.24</sup> <sub>-0.28</sub>	1.47 <sup>+0.11</sup> <sub>-0.10</sub>	26.9 <sup>+1.1</sup> <sub>-0.7</sub>
Fe xvii	15.014	1.94 <sup>+0.32</sup> <sub>-0.33</sub>	1.55 <sup>+0.13</sup> <sub>-0.12</sub>	52.4 <sup>+2.5</sup> <sub>-1.6</sub>
Fe xvii	16.780	2.86 <sup>+0.38</sup> <sub>-0.71</sub>	1.01 <sup>+0.61</sup> <sub>-0.00</sub>	23.1 <sup>+1.9</sup> <sub>-1.2</sub>
Fe xvii <sup>c</sup>	17.051, 17.096	2.52 <sup>+0.70</sup> <sub>-0.64</sub>	1.47 <sup>+0.35</sup> <sub>-0.46</sub>	32.7 <sup>+0.9</sup> <sub>-1.1</sub>
O viii	18.969	3.02 <sup>+0.52</sup> <sub>-0.57</sub>	1.18 <sup>+0.41</sup> <sub>-0.17</sub>	37.0 <sup>+2.8</sup> <sub>-2.6</sub>
N vii	20.9099	4.26 <sup>+2.28</sup> <sub>-1.71</sub>	1.88 <sup>+0.87</sup> <sub>-0.87</sub>	14.8 <sup>+2.3</sup> <sub>-1.9</sub>
O vii	21.602, 21.804	1.62 <sup>+1.33</sup> <sub>-0.79</sub>	2.53 <sup>+0.85</sup> <sub>-0.50</sub>	59.9 <sup>+4.9</sup> <sub>-5.4</sub>

<sup>a</sup> Closely spaced doublets in the Lyman series lines and He-like intercombination lines are fit with a single profile model centered at the emissivity-weighted wavelength of the two components.

<sup>b</sup> For the blended lines fit simultaneously, including the He-like complexes, the total normalization of all the lines in the complex is indicated.

<sup>c</sup> We fit these two blended lines simultaneously, with a fixed normalization ratio of 0.9. Both line profile components were forced to have the same  $\tau_*$  and  $R_o$  values.

Values of the power-law index,  $q$ , that differ significantly from zero (no radial variation) can cause changes in the line profiles that are not insignificant, effectively weighting the emission from parts of the wind according to their velocity (via the beta-velocity law relationship between velocity and radius). However, we find that when we allow  $q$  to be a free parameter the best-fit value is generally very close to zero. For the representative 15.014 line, it is  $q = -0.09$ , and  $q = 0$  is included in the 68 percent confidence region. The general result is consistent with that found for this and other stars (Kramer et al. 2003; Cohen et al. 2006). Thus, to keep the number of free parameters manageable, we fix  $q = 0$ .

The factors discussed above have a very minor influence on the results of the line fitting. However, the remaining factors can have a significant effect.

The velocity-law exponent,  $\beta$ , affects line profiles for two reasons: (1) the velocity law affects the mapping between radius and Doppler shifted wavelength, and so affects the emission profile; and (2) via the continuity equation, it affects the density and so affects both the emission and the absorption. Indeed, for our representative emission line, when we change the value of  $\beta$  from 1 to 0.8, both  $\tau_*$  and  $R_o$  change by 10 to 20 percent. The determinations of  $\beta$  for  $\zeta$  Pup vary from at least 0.9 to 1.15, and so using a value of  $\beta = 1$  seems reasonable, especially as it speeds the calculation of the line profile model by allowing the optical depth integral to be done analytically, so we use that value for all the model fitting results reported here. If, in the future, a new and more accurate determination of  $\beta$  is

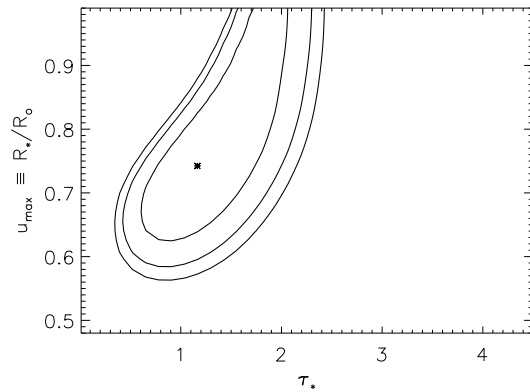
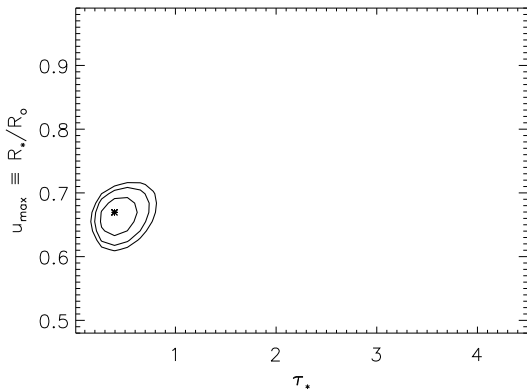
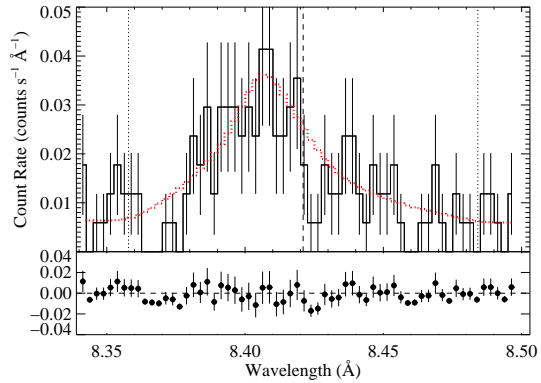
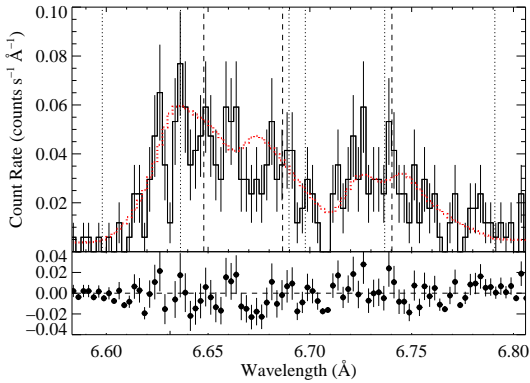
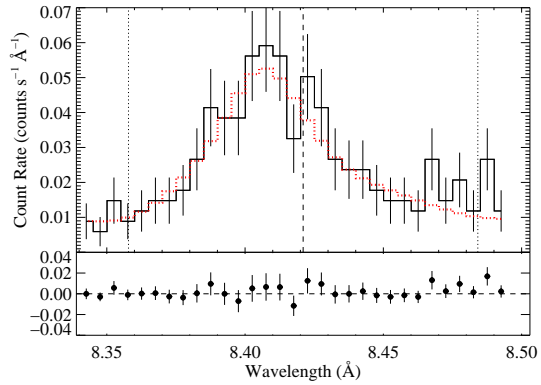
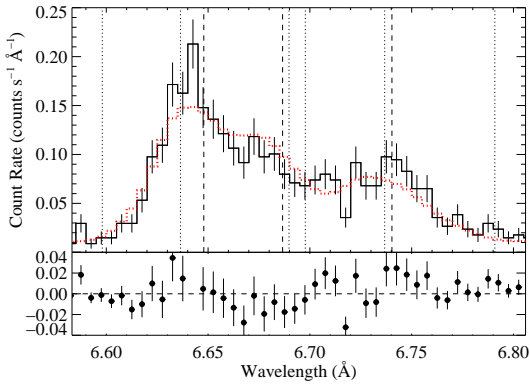
made, and it differs significantly from  $\beta = 1$ , then the results reported in this paper can be scaled accordingly<sup>3</sup>. We also note that the X-ray emitting plasma and the bulk wind that attenuates the X-rays may not necessarily be described by the same beta velocity law. However, there is no independent evidence for this, and with the short post-shock cooling lengths expected in the relatively dense wind of  $\zeta$  Pup, the X-ray emitting plasma in the wind is more likely to have a velocity close to the ambient wind velocity<sup>4</sup>. And furthermore, the observed X-ray emission line widths in  $\zeta$  Pup and other early O supergiants are completely consistent with the  $\beta$  and  $v_\infty$  values inferred from UV and optical spectroscopy of these stars.

The terminal velocity of  $\zeta$  Pup is relatively well established, with reasonable estimates from several different groups that vary by about  $\pm 10$  percent about our adopted value of 2250 km s<sup>-1</sup>. However, when we explored the effect of varying the terminal velocity in our fitting of non-porous

<sup>3</sup> Lowering  $\beta$  from 1 to 0.8 causes the best-fit optical depth of the Fe xvii line at 15.014 Å to go from  $\tau_* = 1.98$  to  $\tau_* = 1.66$ . If the value of  $\beta$  were to be revised upward by a similar amount, the values we derive for  $\tau_*$  from the line profile fitting would have to be revised upward by about 15 percent. The quality of the fits with the different values of  $\beta$  do not differ significantly.

<sup>4</sup> X-ray emitting plasma is too highly ionized to be effectively driven by the photospheric UV radiation field. However, for small enough parcels, the ram pressure of the surrounding wind should keep the post-shock, hot plasma moving at the ambient velocity.

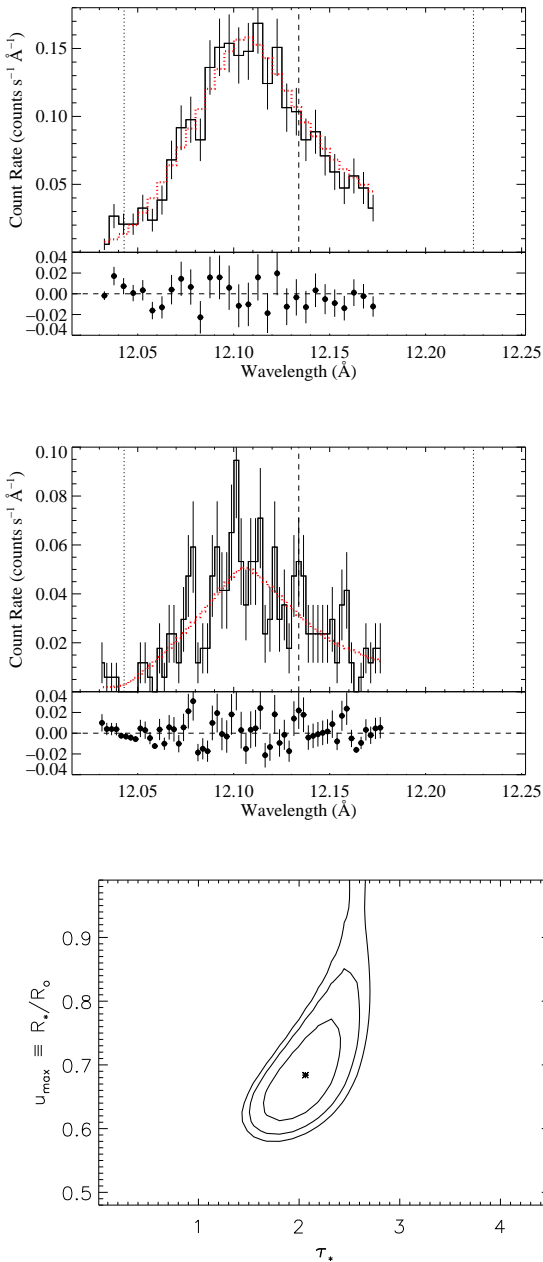




**Figure 6.** The MEG (top) and HEG (middle) measurements of the Si XIII helium-like complex near 6.7 Å, along with the best-fit model. This line complex shows a relatively small degree of blue shift and asymmetry, indicative of a low  $\tau_*$  value, as is expected at short wavelengths, where the wind opacity is smaller. Note that there is a separate set of vertical lines – denoting the rest wavelength and the Doppler shifts associated with the wind terminal velocity – for each of the three components of the line complex (resonance, intercombination, and forbidden lines, from blue to red). We also show the 68, 90, and 95 percent confidence limits in  $\tau_*$ ,  $u_o$  parameter space (bottom).

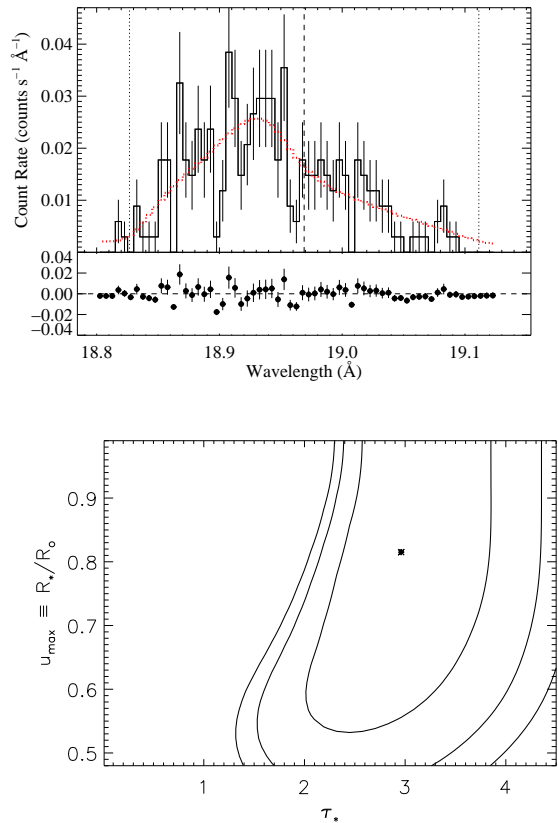
**Figure 7.** The MEG (top) and HEG (middle) measurements of the Mg XII Ly $\alpha$  line at 8.421 Å, along with the best-fit model. The derived value of  $\tau_*$  is significantly higher than that found for the shorter wavelength Si XIII complex shown in the previous figure. We also show the 68, 90, and 95 percent confidence limits in  $\tau_*$ ,  $u_o$  parameter space (bottom).

wind profile models to the 15.014 Å line, we found that the value of  $\tau_*$  was quite sensitive to the assumed wind terminal velocity, even within this relatively narrow range. This is because the blue shift of the line centroid in the dimensionless, scaled wavelength parameter,  $x \equiv (\lambda/\lambda_o - 1)c/v_\infty$ , depends directly on the degree of wind absorption. The same observed profile appears more blue shifted in scaled wavelength units if the terminal velocity is (assumed to be) smaller. Our tests with the 15.014 Å line show that the best-fit value for  $\tau_*$



**Figure 8.** The MEG (top) and HEG (middle) measurements of the Ne x Ly $\alpha$  line at 12.134 Å, along with the best-fit model. This line shows an intermediate degree of blue shift and asymmetry, indicative of an intermediate  $\tau_*$  value, as is expected at its wavelength, where the wind opacity is larger than at the wavelength of the Mg XII Ly $\alpha$  line, but not as large as at longer wavelengths. We also show the 68, 90, and 95 percent confidence limits in  $\tau_*$ ,  $u_o$  parameter space (bottom).

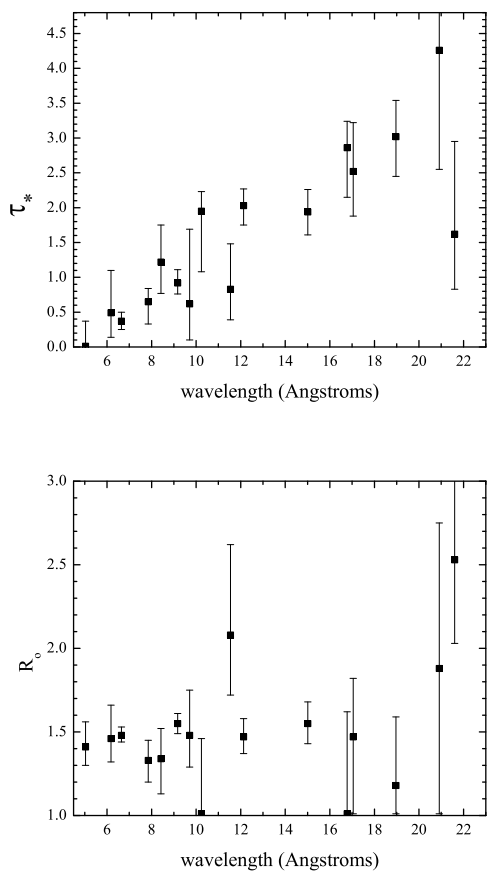
ranges from 2.16 to 1.35 when we use terminal velocities between 2200 km s $^{-1}$  and 2485 km s $^{-1}$ . This variation is larger than that caused by every other parameter uncertainty and assumption we have explored. Thus, while we consider the value of  $v_\infty = 2250$  km s $^{-1}$  to be quite reliable, future reassessments of this parameter will necessitate a rescaling of



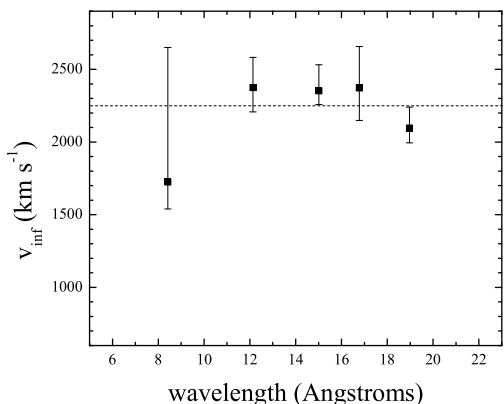
**Figure 9.** The MEG (top) measurements of the O VIII Ly $\alpha$  line at 18.969 Å, along with the best-fit model. This line shows a relatively large degree of blue shift and asymmetry, indicative of a higher  $\tau_*$  value, as is expected at longer wavelengths, where the wind opacity is larger. We did not include the very weak HEG data in the analysis of this line. We also show the 68, 90, and 95 percent confidence limits in  $\tau_*$ ,  $u_o$  parameter space (bottom).

the optical depth, and mass-loss rate, results we report in this paper.

As a final test, we can treat the terminal velocity as a free parameter of the model. This enables us to see what value of the terminal velocity is preferred by the X-ray spectral data themselves. In general, the constraints on  $v_\infty$ , while letting the other model parameters be free to vary, were not strong. But for the highest signal-to-noise lines in the spectrum, relatively tight constraints could be derived. We show the results for fitting the five most useful lines in Fig. 11. As the figure shows, these lines are all consistent with our adopted value of  $v_\infty = 2250$  km s $^{-1}$ . This, of course, gives us added confidence that the value we use for the model fitting is reasonable. And, in fact, the small error bars on most of these determinations also show that significantly smaller and larger values are ruled out. The kinematics of the hot, X-ray emitting plasma seem to be the same as that of the bulk wind.



**Figure 10.** Values of  $\tau_*$  (top) and  $R_o$  (bottom) derived from the model fits, shown with their 68 percent confidence limits. Line complexes and blends that were fit with multiple model components are represented by only one point.



**Figure 11.** Values of the terminal velocity derived from fitting five strong lines with a wind profile model for which  $v_\infty$  was allowed to be a free parameter (along with  $\tau_*$ ,  $R_o$ , and the normalization. The bulk wind terminal velocity adopted from the analysis of UV profiles is indicated by the horizontal line. *This figure will be remade with IDL, for consistency in the axis labeling, etc.*

## 5 DISCUSSION

The most obvious new and significant result of the profile model fitting is the wavelength trend in the derived values of the fiducial optical depth,  $\tau_*$ , shown in the top panel of Fig. 10. The value of this parameter, which is proportional to both the mass-loss rate and the opacity of the bulk wind, increases with wavelength, which is exactly what is expected from the form of the atomic opacity. The null hypothesis of a constant value of  $\tau_*$  is rejected with greater than 99.9 percent confidence ( $\chi^2_\nu = 5.4$  for 15 degrees of freedom). We therefore fit a model of wavelength-dependent  $\tau_*$ , in which the wavelength dependence derives entirely from the atomic opacity,  $\kappa(\lambda)$ .

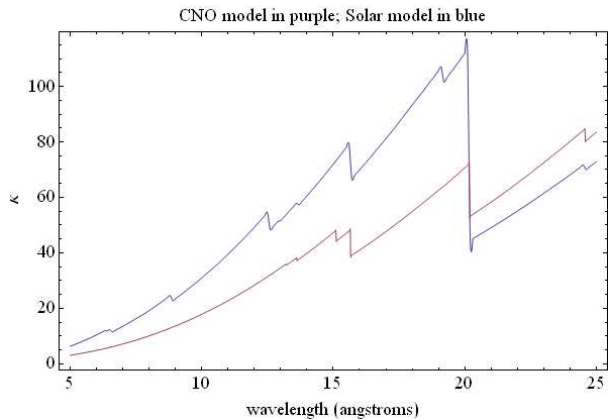
While it may seem obvious that there should be a trend in the fiducial optical depth with wavelength, this result is quite significant, in that the assumed lack of such a trend is the basis for claims that large-scale clumping and the associated wind porosity are the cause of the modest profile blue shifts and asymmetry (Osino et al. 2006). In the following subsections, we show how a realistic wind opacity model naturally explains the observed wavelength trend, and then how such a model can be used to make a quantitative determination of the mass-loss rate of  $\zeta$  Pup.

### 5.1 The opacity model

The opacity model depends on the abundances and, to a lesser extent, the ionization balance of the bulk stellar wind (i.e. the cooler, unshocked component). The dominant source of opacity is photoelectric absorption, from the K-shell of abundant elements between N and Si, and also the L-shell of Fe. We have computed a wind opacity model using CMFGEN. The model is constrained by UV and optical spectra, so the wind ionization balance and abundances are consistent with observations. Specifically, the model has  $Y_{\text{He}} = 0.16$ ,  $(Z/Z_\odot)_\text{C} = 0.03$ ,  $(Z/Z_\odot)_\text{N} = 3.35$ ,  $(Z/Z_\odot)_\text{O} = 0.14$ , and  $(Z/Z_\odot)_\text{Fe} = 0.84$ , where the reference solar abundances are taken from Grevesse & Sauval (1998). We show this wind opacity model, at a single radius ( $r = 2 R_*$ ) in Fig. 12, along with a solar-abundance model. The opacity is lower at most wavelengths in the custom-computed model, primarily because the total abundance of metals (and most crucially the sum of nitrogen and oxygen) is subsolar.

Using either of these models of the opacity, and values for the stellar radius and wind terminal velocity from Table 1, we can construct a wavelength-dependent model of  $\tau_*$ , for which the mass-loss rate is the only free parameter. Fits with both the custom-computed wind opacity model and the solar abundance model are equally good ( $\chi^2_\nu \approx 0.8$ ), although a higher mass-loss rate of  $\dot{M} = 3.1 \times 10^{-6} M_\odot \text{ yr}^{-1}$  is found with the cmfgen model, due to its lower overall opacity. The solar abundance opacity model, which should provide a lower limiting case, gives  $\dot{M} = 1.9 \times 10^{-6} M_\odot \text{ yr}^{-1}$ . The formal uncertainties on these derived mass-loss rates, due solely to the finite error bars on the individual  $\tau_*$  determinations are about 10 percent.

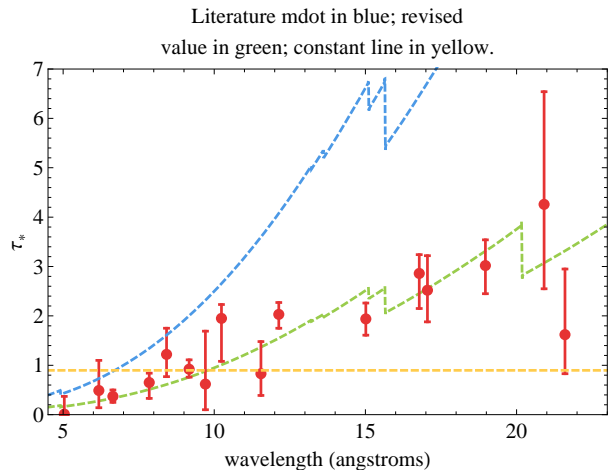
The best-fit  $\tau_*$  model, using the custom-computed opacities, is shown in Fig. 13, along with the  $\tau_*$  model computed using the standard unclumped mass-loss rate from the literature,  $\dot{M} = 8.3 \times 10^{-6} M_\odot \text{ yr}^{-1}$ . The best-fit mass-loss rate



**Figure 12.** The wavelength dependent opacity of the wind of  $\zeta$  Pup computed with CMFGEN (purple), along with a solar-abundance opacity model (blue). Note the prominent K-shell edge of oxygen near 20 Å in the solar abundance model. In the custom-computed model, this decrement is much more modest, due to the underabundance of O and overabundance of N. The overall reduction in the opacity at most wavelengths in the custom-computed model is the result of overall subsolar abundance in the CMFGEN model. *We will remake this figure in b/w without a title above the plot. I thought we'd wait till we had the final CMFGEN model, including the relevant metals.*

is close to a factor of three lower. If solar abundances are assumed for the opacities, the factor is closer to four. The two models are compared in Fig. 14, and have a very similar shape, implying that even with better data it would be difficult to distinguish them based on the X-ray data alone.

Taking a closer look at the atomic opacity, we can see in Fig. 15 that the most leverage regarding the wavelength dependence of the opacity, and hence of  $\tau_*$ , comes at the shortest wavelengths, well below the Fe L-shell edges near 15 Å. The successive K-shell edges of N and O at longer wavelengths, conspire to make the opacity rather flatter in that region than the generally expected  $\kappa \propto \lambda^3$  relationship seen from individual elements' photoionization cross sections. Furthermore, the paucity of useful emission lines longward of the O K-shell edge makes it difficult to discriminate among various wind opacity models, although in principle, lines longward of this edge could enable us to diagnose the altered CNO-processed abundances with some certainty. And emission lines longward of the N K-shell edge near 26 Å would be especially useful, but there are none in the *Chandra* spectrum. The N VII Ly $\beta$  line at 20.910 Å is quite weak and does not provide a strong constraint on  $\tau_*$ , although it does favor the custom-computed (CMFGEN) opacity model. The longest wavelength line which we are able to reliably fit is the helium-like O VII complex near 21.8 Å. We fit the resonance and intercombination lines simultaneously (the forbidden line is not present due to  $2^3S - 2^3P$  photoexcitation by the photospheric UV field), with the profile parameters  $\tau_*$  and  $R_o$  tied together for the two lines. However, the resonance line in this complex may be subject to resonance scattering (Leutenegger et al. 2007) – it may be optically thick to its own radiation (as distinct from the effects of continuum opacity of the overlying wind that leads to the observed skewness and blue shifts in all of the line profiles). Resonance scattering tends to make broadened, asymmet-



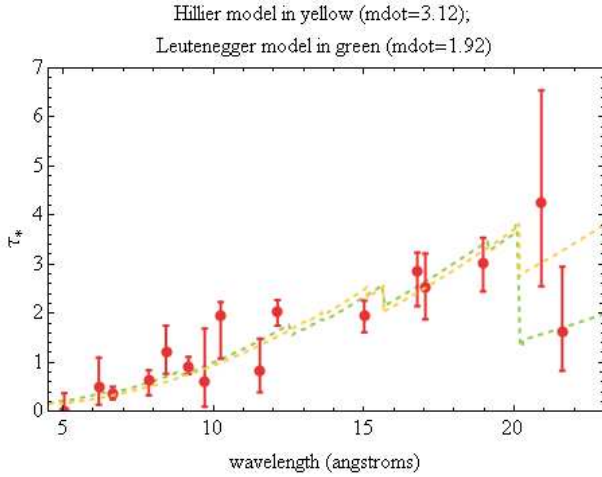
**Figure 13.** Values of  $\tau_*$  derived from the line-profile model fits, shown as points with error bars (same as the top panel of Fig. 10). The value of  $\tau_*$  expected from the literature mass-loss rate of  $8.3 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$  is shown as the upper solid (blue) line. This model has the same wavelength dependence as the non-solar, custom-computed opacity shown in Fig. 12 (purple line). Treating the mass-loss rate as a free parameter, the best fit value of  $3.1 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$  is shown as the lower, varying solid (green) line. This model provides a formally good fit. The horizontal dashed (yellow) line is the best-fit constant  $\tau_*$  model, as would be expected for a porosity-dominated wind. It does not provide a good fit to the data. *This figure will be remade in b/w, with line styles that are consistent with those shown in the related figures. The title at the top will also be removed. Again, I am going to wait to do this until the CMFGEN opacity model is finalized.*

ric, and blue shifted lines more symmetric, and thus the  $\tau_*$  value we derive from fitting this complex may be somewhat underestimated. If this is the case, then this line complex too would favor the custom-computed, subsolar abundance wind opacity model, as shown in Fig. 14. We also note that the only other line in the spectrum that is likely to be optically thick to resonance scattering is the O VII Ly $\alpha$  line at 18.969 Å, so the  $\tau_*$  determination for that line may also be somewhat underestimated.

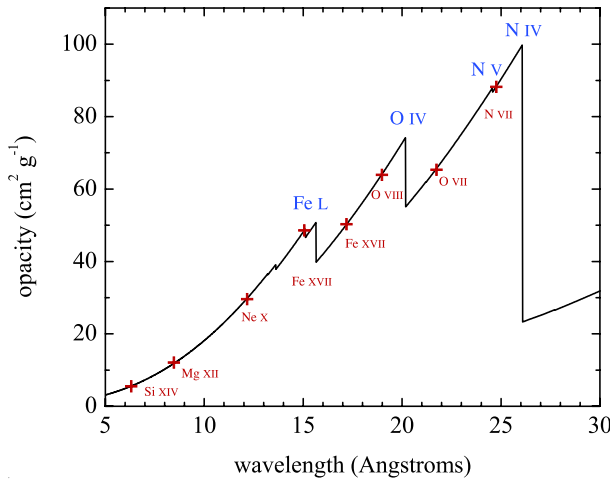
We also can see from a careful inspection of the opacity model that the mass-loss rate determination from fitting a set of  $\tau_*$  values is mostly sensitive to the cross section contributions from N, O, and Fe. Alterations of O and N abundances due to CNO processing will have only a modest effect on the results, however. The sum of the contributions of C, N, and O (as well as He) is what affects the overall opacity level longward of about 15 Å, with Fe making a significant contribution at shorter wavelengths. This demonstrates that accurate determinations of abundances for O stars are perhaps the biggest factor in enabling the determination of clumping-independent mass-loss rates from high-resolution X-ray spectra. But when fitting a large ensemble of lines that span a relatively wide range of wavelengths, knowing the overall metallicity is probably sufficient.

## 5.2 Rejection of Gray Opacity and the Porosity Model

The rejection of the hypothesis of a gray effective opacity (wavelength-independent  $\tau_*$ ) argues strongly against the



**Figure 14.** The best-fit model, with  $\dot{M} = 3.1 \times 10^{-6}$ , shown in Fig. 13, is shown here in yellow, while the best-fit solar abundance  $\tau_*$  model is shown in green. The fits are of equally good quality, while the solar abundance model has a lower mass-loss rate. *Here too, we will remake the figure in b/w, without a title at the top, and using line styles that are consistent with the previous figures. Note that the colors here don't match those in the previous figure.*



**Figure 15.** The same non-solar, custom-computed (CMFGEN) wind opacity model shown (as the purple line) in Fig. 12, but with the major ionization edges labeled (above the opacity trace), and the wavelengths of some strong lines indicated (by crosses, with labels below the opacity trace). *Again, this figure will be remade in IDL, for consistency in typeface and axis labeling; it will be fully b/w. I'm a little concerned that this figure will be confusing because of the different x-axis range (as the previous couple of figures). But I want to show the N edge.*

dominance of large-scale clumping and its associated wind porosity in setting the observed properties of O star X-ray spectra. It has been shown that very large-scale clumping, in which individual clumps are optically thick in the soft X-ray continuum, can lead to a reduction in the effective opacity that might explain the only modestly shifted and asymmetric profiles observed in many O star X-ray spectra (Oskinova et al. 2006; Owocki & Cohen 2006). In this scenario, it is the physical cross section of the clumps, and not the atomic cross sections, that governs the wind opacity.

Thus the expectation is that a porosity dominated wind would lead to X-ray emission line profiles with similar morphologies and no wavelength dependence across the observed spectral range. This scenario requires a very large clump scale, however, with porosity lengths in excess of a stellar radius (Oskinova et al. 2006; Owocki & Cohen 2006), where the porosity length represents the interclump mean free path in the limit of very optically thick clumps. Furthermore, preliminary fits of X-ray line profile models that explicitly include porosity as a free parameter do not, in fact, provide better fits to the observed profiles than models that simply have reduced mass-loss rates (Cohen et al. 2008). In an upcoming paper, we will show that this conclusion holds for all strong lines in the *Chandra* spectrum of  $\zeta$  Pup, and furthermore, that to reproduce the observed profiles with the standard, unclumped mass-loss rate of  $\dot{M} = 8.3 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$ , requires porosity lengths of roughly  $3 R_*$  (which we have already shown is the case for the high signal-to-noise Fe XVII line at 15.014 Å in Cohen et al. (2008)).

### 5.3 Sources of Uncertainty in the Mass-Loss Rate Determination

The significant sources of uncertainty in the mass-loss rate determination we have derived from the fits to the ensemble of  $\tau_*$  values derived from fitting the individual line profiles come from three sources. The first is the formal uncertainty on the mass-loss rate model that stems from the uncertainties on the individual line profile fits (represented by the error bars on the  $\tau_*$  points in Fig. 13, for example). For the custom-computed CMFGEN opacity model, the 68 percent confidence limit range on the fitted mass-loss rate extends from 2.89 to  $3.35 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$ , representing an uncertainty of a little less than 10 percent.

The second source of uncertainty arises from our imperfect knowledge of the wind terminal velocity (and, most importantly, the terminal velocity of the X-ray emitting plasma itself). However, as we have shown (see Fig. 11), the data themselves indicate that our adopted terminal velocity of  $v_{\infty} = 2250 \text{ km s}^{-1}$  is well supported. Three of the lines we show in that figure have best-fit terminal velocity values near  $2350 \text{ km s}^{-1}$ , which is also the terminal velocity derived from a careful analysis of the UV line profiles by Haser (*ref: 1997 dissertation*). When we refit the representative Fe XVII line at 15.014 Å using this higher terminal velocity, we found a reduction in our derived  $\tau_*$  value of 15 percent. If this scaling holds for all lines, then using this slightly higher value of the terminal velocity will lead to a downward revision of our derived mass-loss rate of about 15 percent. (Note that the terminal velocity enters into the denominator of the expression for  $\tau_*$ , and that will mitigate this adjustment slightly.)

Finally, we estimate that the abundances derived for  $\zeta$  Pup from the CMFGEN modeling and constraints from the UV and optical data have a precision of about 30 percent. We can see from the comparison of the CMFGEN model to the solar abundance model that the mass-loss rate varies by about 30 percent between these two assumed opacity models, although the solar abundance model is included in our analysis not so much as a realistic alternate model, but simply as a plausible upper bound to the atomic opacity; the custom-computed CMFGEN model is more realistic due

to the constraints on it provided by observations in other wavelength bands. Thus, a conservative estimate of the allowed range of the mass-loss rate of  $\zeta$  Pup derived from the X-ray line profile fitting is roughly 2 to  $4 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$ , with our best estimate being  $3.1 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$ . This mass-loss rate is only a little lower than the minimum clumping mass-loss rate of  $4.2 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$  (Puls et al. 2006), implying a small amount of clumping in the outer wind, and a small adjustment to the clumping factor in the inner wind determined by Puls et al. (2006).

#### 5.4 Location of the X-ray Emitting Plasma

The analysis of the sixteen lines and line complexes in the *Chandra* spectrum of  $\zeta$  Pup also enables us to derive values of the onset radius of the wind-shock X-ray emission from the profiles. These results are shown in the lower panel of Fig. 10, and are completely consistent with the expectations of the wind-shock structure induced by the line-driven instability (Feldmeier et al. 1997; Runacres & Owocki 2002). That is, an onset radius of  $R_o \approx 1.5 R_*$ . We have searched for a trend with wavelength in these values and found none (an unweighted fit of a linear trend shows a modest increase with wavelength, but that result is significant at only the one sigma level, and when we perform a weighted fit – with the weights inversely proportional to the uncertainties on the individual measurements – the significance is less than one sigma). Thus, the simplest interpretation is that there is a universal radius of the onset of X-ray emission and it occurs near  $1.5 R_*$  (half a stellar radius above the photosphere). This result had already been noted by Kramer et al. (2003), though we show it more robustly here. This same result can also be seen in the late O supergiant  $\zeta$  Ori (Cohen et al. 2006). And this result is also consistent with the joint analysis of X-ray line profile shapes and helium-like forbidden-to-intercombination line ratios for four O stars as described by Leutenegger et al. (2006).

#### 5.5 Comparison with Previous Analyses

Finally, let us consider why we have found a trend in wavelength for the fiducial optical depth values,  $\tau_*$ , derived from the same *Chandra* data that led Kramer et al. (2003) to report that there was no obvious trend. The two biggest factors that have led to this new result are our more careful assessment of line blends and our inclusion of many weak, but important, lines at short wavelength. Kramer et al. (2003) included only one line shortward of the Ne x Ly $\alpha$  line at 12.134 Å, whereas we report on nine lines or line complexes in this range (including two helium-like complexes, which Kramer et al. (2003) excluded from their analysis). While many of these lines are weak and do not provide very strong constraints when considered individually, taken together, they do provide strong constraints. As far as line blends are concerned, Kramer et al. (2003) included the N VII Ly $\alpha$  line at 24.781 Å and the Fe XVII complex near 15.26 Å, both of which we have determined are too blended to enable the extraction of reliable information about their intrinsic profile shapes. Furthermore, we properly account for the blended Fe XVII lines at 17.051 and 17.096 Å, fitting them simultaneously, while Kramer et al. (2003) fit them as a single line.

Our use of detailed models of the effective area and resolution also may play a small role in making our results more reliable than those in the initial paper.

## 6 CONCLUSIONS

By quantitatively analyzing all the X-ray line profiles in the *Chandra* spectrum, we have determined a mass-loss rate of  $3 \pm 1 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$ . The largest uncertainty arises from the abundances in the atomic opacity model. This method of mass-loss rate determination from X-ray profiles is a potentially powerful tool for addressing the important issue of the actual mass-loss rates of O stars. Care must be taken in the profile analysis, however, as well as in the interpretation of the trends found in the derived  $\tau_*$  values. It is especially important to use a realistic model of the wind opacity. And for O stars with weaker winds, especially, it will be important to verify that the X-ray profiles are consistent with the overall paradigm of embedded wind shocks. Here, an independent determination of the terminal velocity of the X-ray emitting plasma by analyzing the widths and profiles of the observed X-ray lines themselves will be crucial. In the case of  $\zeta$  Pup, we have shown that the X-ray profiles are in fact consistent with the same wind kinematics seen in UV absorption line spectra of the bulk wind. And the profile analysis also strongly constrains the onset radius of X-ray production to be about  $r = 1.5 R_*$ .

A major conclusion of from the profile analysis is that there is no need to invoke large scale porosity to explain individual line profiles, as the overall wavelength trend argues strongly against porosity as being the dominant cause of the reduced effective opacity of the wind. Rather, the opacity is completely consistent with the expected wavelength-dependent atomic opacity, with the lower-than-expected wind optical depths being due, simply, to a reduction in the wind mass-loss rates. This reduction (a factor of three) is fully consistent with other recent determinations that correct the traditional wind mass-loss rate diagnostics for small-scale clumping that affects density-squared diagnostics and ionization corrections.

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*Coauthors: any acknowledgments you'd like to add?*

## REFERENCES

- Berghoefter T. W., Baade D., Schmitt J. H. M. M., Kudritzki R.-P., Hillier D. J., Pauldrach A. W. A., 1996, *A&A*, 306, 899
- Bouret J.-C., Lanz T., Hillier D. J., 2005, *A&A*, 438, 301
- Cash W., 1979, *ApJ*, 228, 939
- Cassinelli J.P., Miller N.A., Waldron W.L., MacFarlane J.J., Cohen D.H., 2001, *ApJ*, 554, L55
- Castor J. I., Abbott D. C., Klein R. I., 1975, *ApJ*, 195, 157
- Cohen D.H., Leutenegger M.A., Grizzard K.T., Reed C.L., Kramer R.H., Owocki S.P., 2006, *MNRAS*, 368, 1905

- Cohen D. H., Leutenegger M. A., Townsend R. H. D., 2008, in *“Clumping in Hot Star Winds”*, eds Hamann W.-R., Feldmeier A., and Oskinova L. M., (Universitätsverlag, Potsdam), p209
- Dessart L., Owocki S. P., 2003, A&A, 406, L1
- Feldmeier A., Oskinova L., Hamann W.-R., 2003, A&A, 403, 217
- Feldmeier A., Puls J., Pauldrach A.W.A., 1997, A&A, 322, 878
- Fullerton A. W., Massa D. L., Prinja R. K., 2006, ApJ, 637, 1025
- Glebocki R., Gnacinski P., Stawikowski A., 2000, Acta Astron., 50, 509
- Grevesse N., Sauval A. J., 1998, Sp. Sc. Rev., 85, 161
- Hillier D. J., Miller D. L., 1999, ApJ, 519, 354
- Hillier D. J., Kudritzki R.-P., Pauldrach A. W. A., Baade D., Cassinelli J. P., Puls J., Schmitt J. H. M. M., 1993, ApJ, 276, 117
- Ignace R., 2001, ApJ, 549, 119
- Kahn S.M., Leutenegger M.A., Cottam J., Rauw G., Vreux J.-M., den Boggende A.J.F., Mewe R., Güdel M., 2001, A&A, 365, L312
- Kramer R. H., Cohen D. H., Owocki S. P., 2003, ApJ, 592, 532
- Lamers H. J. G. L. M., Leitherer C., 1993, ApJ, 412, 771
- Leutenegger M. A., Paerels F. B. S., Kahn S. M., Cohen D. H., 2006, ApJ, 650, 1096
- Leutenegger M. A., Owocki S. P., Paerels F. B. S., Kahn S. M., 2007, ApJ, 659, 642
- Lucy L. B., White R. L., 1980, ApJ, 241, 300
- MacFarlane J. J., Cassinelli J. P., Welsh B. Y., Vedder P. W., Vallergera J. V., Waldron W. L., 1991, ApJ, 380, 564
- Marshall H.L., Dewey D., Ishibashi K., 2004, SPIE, 5165, 457
- Oskinova L., Feldmeier A., Hamann W.-R., 2004, A&A, 422, 675
- Oskinova L., Feldmeier A., Hamann W.-R., 2006, MNRAS, 372, 313
- Owocki S.P., Castor J.I., Rybicki G.B., 1988, ApJ, 335, 914
- Owocki S.P., Cohen D.H., 2001, ApJ, 559, 1108
- Owocki S.P., Cohen D.H., 2006, ApJ, 648, 565
- Pauldrach A. W. A., Puls J., Kudritzki R.-P., 1986, A&A, 164, 86
- Perryman M.A.C., et al., 1997, A&A, 323, L49
- Hamann W.-R., Feldmeier A., Oskinova L. M., 2008, *Clumping in Hot Star Winds*. Potsdam, Universitätsverlag
- Press W.H., Flannery B.P., Teukolsky S.A., Vetterling W.T., 2007, *Numerical Recipes*, 3<sup>rd</sup> edition. Cambridge, Cambridge University Press
- Prinja R. K., Barlow M. J., Howarth I. D., 1990, ApJ, 361, 607
- Puls J., Kudritzki R.-P., Herrero A., Pauldrach A. W. A., Haser S. M., Lennon D. J., Gabler R., Voels S. A., Vilchez J. M., Wachter S., Feldmeier A., 1996, A&A, 305, 171
- Puls J., Markova N., Scuderi S., Stanghellini C., Taranova O. G., Burnley A. W., Howarth I. D., 2006, A&A, 454, 625
- Repolust T., Puls J., Herrero A., 2004, A&A, 415, 349
- Runacres M. C., Owocki S.P., 2002, A&A, 381, 1015
- van Leeuwen R., 2007a, A&A, 474, 653
- van Leeuwen R., 2007b, *Hipparcos, the New Reduction of the Raw Data*. Springer, Berlin
- Vanbeveren D. de Loore C., van Rensbergen V., 1998, A&A Rev., 9, 63