

# Measuring the shock heating rate in the winds of O stars using X-ray line spectra

David H. Cohen,<sup>1\*</sup> Zequn Li,<sup>1</sup> Kenneth G. Gayley,<sup>2</sup> Stanley P. Owocki,<sup>3</sup>  
Jon O. Sundqvist,<sup>3,4</sup> Véronique Petit,<sup>3</sup> Maurice A. Leutenegger<sup>5,6</sup>

<sup>1</sup>*Swarthmore College, Department of Physics and Astronomy, Swarthmore, Pennsylvania 19081, USA*

<sup>2</sup>*University of Iowa, Department of Physics and Astronomy, Iowa City, Iowa 52242, USA*

<sup>3</sup>*University of Delaware, Bartol Research Institute, Newark, Delaware 19716, USA*

<sup>4</sup>*Institut für Astronomie und Astrophysik der Universität München, Scheinerstr. 1, 81679 München, Germany*

<sup>5</sup>*NASA/Goddard Space Flight Center, Code 662, Greenbelt, Maryland 20771, USA*

<sup>6</sup>*CRESST and University of Maryland, Baltimore County, MD 21250, USA*

4 July 2014

## ABSTRACT

We present a new method for using measured X-ray emission line fluxes from O stars to determine the shock heating rate due to instabilities in their radiation-driven winds. The high densities of these winds means that their embedded shocks quickly cool by local radiative emission, while cooling by expansion should be negligible. Ignoring for simplicity any non-radiative mixing or conductive cooling, the method presented here exploits the idea that the cooling post-shock plasma systematically passes through the temperature characteristic of distinct emission lines in the X-ray spectrum. In this way, the observed flux distribution among these X-ray lines can be used to construct the cumulative probability distribution – or cumulative initial-temperature distribution (Gayley 2014) – of shock strengths that a typical wind parcel encounters as it advects through the wind. Determining this shock-strength distribution requires applying a wind absorption correction to the flux from each observed emission line, using results from a well-developed line profile fitting model (Cohen et al. 2014). We apply this new method to *Chandra* grating spectra from five O stars with X-ray emission indicative of embedded wind shocks in effectively single massive stars. The results for all the stars are quite similar: the average wind mass element passes through roughly one shock that heats it to at least  $10^6$  K as it advects through the wind, and the *cumulative* distribution of shock strengths is a strongly decreasing function of temperature, consistent with a negative power-law of index  $n \approx 3$ , implying a *marginal* distribution of shock strengths that scales as  $T^{-4}$ , and with hints of an even steeper decline or cut-off above  $10^7$  K.

**Key words:** stars: early-type – stars: winds, outflows – X-rays: stars

## 1 INTRODUCTION

Embedded Wind Shocks (EWS) are the source of the ubiquitous soft X-ray emission seen in O stars. This is confirmed by the significantly Doppler-broadened X-ray emission lines observed with *Chandra* and *XMM-Newton* (Kahn et al. 2001; Cassinelli et al. 2001; Kramer et al. 2003). The EWS are generally thought to be associated with the Line Driven Instability (LDI) that is intrinsic to any radiation-driven flow in which spectral lines mediate the transfer of momentum from the radiation field to matter (Milne 1926;

Lucy & Solomon 1970; Owocki et al. 1988). Indeed, hydrodynamics simulations show numerous shocks and associated clumped structure in O-stars winds (Cooper 1994; Feldmeier et al. 1997; Runacres & Owocki 2002). Modeling indicates that the extent to which the instability is seeded by photospheric variability and limb darkening (Feldmeier et al. 1997; Sundqvist & Owocki 2013) can have a strong effect on the shock structure and X-ray emission that is produced. Additionally, multidimensional effects (Dessart & Owocki 2003, 2005) could also have a significant effect on the shock heating and X-ray emission that thus far has been explored numerically only in one-dimensional simulations.

\* E-mail: dcohen1@swarthmore.edu

vide constraints on wind models and the physics of the LDI and the associated EWS. However, the X-ray emission levels from embedded wind shocks in massive stars are affected both by the impulsive shock heating and by the cooling of the post-shock plasma, which can be through both radiative and non-radiative channels. The fundamental questions are about the efficiency and physics of the process heating the wind plasma to X-ray emitting temperatures, so it is the nature of the heating that we wish to use X-ray data to constrain.

When multiple cooling channels, including non-radiative ones, contribute significantly, the various cooling processes as well as the heating processes have to be modeled in order to compare theory with observations. In the dense winds of many O stars, however, the cooling by radiation is prompt and occurs locally – the cooling lengths are short. This situation presents us with the opportunity to parametrize the heating directly, which is the approach taken below.

However, a complication for our simple radiatively cooled picture has been recently emphasized by Owocki et al. (2013). Such rapidly cooling, radiative shocks are subject to thin shell instabilities which can lead to mixing-related reductions in X-ray emission. Because of possible mixing effects, the shock heating rates we derive in this paper under the assumption of pure radiative cooling represent lower limits. But to the extent that mixing is negligible, then the fact that radiative cooling is the dominant cooling mechanism enables us to extract shock heating constraints directly from observed X-ray emission line spectra of O stars (Gayley 2014) by exploiting the fact that the shock-heated wind plasma cools locally via the emission of X-rays in different lines, each one with a different temperature-dependent emissivity. It is from this X-ray emission line spectrum that we can reconstruct the shock heating rate and also its temperature dependence.

The X-ray emitting wind plasma in massive stars is usually assumed to be well described by the coronal approximation: statistical equilibrium with collisional ionization from the ground state balanced by radiative and dielectronic recombination, and collisional excitation from the ground state balanced by spontaneous emission. This gives rise to a spectrum dominated by emission lines from modestly excited states to the ground state of highly ionized metals, with a modest contribution from bremsstrahlung and recombination continuum emission. Each emission line has an emissivity that is a relatively peaked, strong function of temperature, following the temperature dependence of the ionization balance and excitation rates. In this way, each line probes a relatively narrow range of plasma temperatures.

The instantaneous X-ray luminosity from a coronal plasma is simply equal to the combined emissivity of all the lines (and continuum processes) multiplied by the emission measure (EM), which is the volume integral of the particle number density squared. This particular dependence arises from the two-body nature of the excitation process of the emission lines (and of the bremsstrahlung and recombination). The temperature distribution of the plasma has thus been traditionally characterized by a differential emission measure (DEM),  $\frac{dEM}{dT}$ , which can be described by a continuous function or by a sum of isothermal components, perhaps taken to approximate a continuous distribution with some

structure. While subject to various data and analysis constraints and ambiguities (see e.g. Gayley 2014), techniques exist for determining a “best-fit” DEM from an observed spectrum (e.g. Kaastra et al. 1996; Liefke et al. 2008). However, as noted above, such a plasma temperature distribution combines both the desired information about shock heating rates and distributions with extraneous and often complex and incomplete information about the cooling history of the hot plasma.

The emission measure is problematic for another reason as well. Namely, the density-squared dependence means that a given mass of heated plasma will have a higher emission measure and radiate faster if it is confined to a smaller volume. Therefore, the differential emission measure of post-shock plasma will depend not only on the heating rate and the cooling rate, but also on the local post-shock density, seemingly causing free parameters of any model that might be fit to data to proliferate. However, a key insight about wind shock X-ray emission that makes the analysis much simpler is that for radiative shocks, the overall X-ray luminosity from an ensemble of wind shocks does not depend on the emission measure behind each shock, but rather on the kinetic energy flux across each shock front (Antokhin et al. 2004; Owocki et al. 2013; Gayley 2014). In the strong shock limit, the kinetic energy added as the wind plasma traverses the shock front is converted to thermal energy. And by definition, if the shock is radiative, that thermal energy will be converted to photons.

A related consequence is that the spectrum of X-ray photons emitted by this radiatively cooled plasma from the time it is impulsively heated until it returns to its pre-shock temperature is independent of how rapidly the plasma cools (Antokhin et al. 2004; Zhekov & Palla 2007). In fact, the luminosity of any particular X-ray emission line is simply given by the energy injected into the plasma by its passage across the shock front multiplied by the *fraction* of the total plasma emissivity that is due to that line (Gayley 2014). We note the similarity to the treatment of X-ray emission from cooling flows in galaxy clusters (Peterson & Fabian 2006).

Furthermore, since plasma impulsively heated to a given temperature as it crosses a shock front will radiate at that temperature and, as time goes on, at every lower temperature until it is fully cooled, an emission line that forms at a given temperature will probe shocks of every higher temperature. And the contribution of a plasma mass parcel to that line will not depend on the heating and cooling history of the parcel, aside from the requirement that the parcel was heated to at least the characteristic temperature of the line. In this way, we can use an ensemble of X-ray emission lines, each with a different temperature dependent emissivity, to derive not only an overall shock heating rate but also to derive the cumulative distribution of shock temperatures. Gayley (2014) has recently presented a comprehensive formalism for analyzing such cumulative initial-temperature distributions for any impulsively heated X-ray emitting plasma within the context of simplified heuristic emissivity functions, and we extend that approach here for detailed emissivities of real lines seen in O-star winds.

In order to obtain the individual X-ray emission line luminosities from the measured line fluxes, a crucial step is to correct not just for the distance to the star, and for the effects of interstellar attenuation, but also for wind attenua-

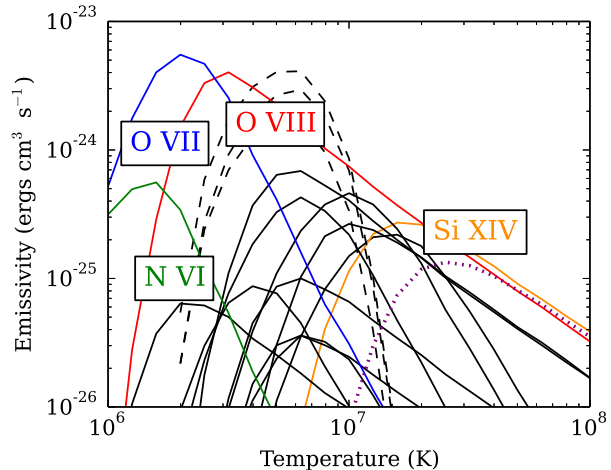
tion of the emitted X-rays. At any given instant only a small fraction of the wind is shock-heated, while most of it is relatively cold and not highly ionized and so it can efficiently absorb X-rays. While accounting for this wind absorption can be done via detailed modeling of the spatial distribution of the X-ray emitting wind plasma and the radiation transport through the absorbing wind simultaneously (e.g. Hervé et al. 2012), the procedure is much more tractable when the wind absorption correction is done separately for each line and independently of the modeling of the heating and cooling. Here we use the wind absorption optical depths derived by Cohen et al. (2014) via X-ray line-profile fitting to make this wind absorption correction.

In this paper we apply the cumulative initial-temperature distribution approach described by Gayley (2014) to the specific case of the dense winds of O stars with radiative wind shocks driving their EWS X-ray emission. Specifically, we analyze the line-rich *Chandra* grating spectra of five OB stars (Cohen et al. 2014) in order to derive the shock heating rates in their winds. In §2 we describe quantitatively how the measured line fluxes are related to the wind shock heating rate and we also describe the data used in the analysis. In §3 we present the results for the program stars. And in §4 we discuss the implications of our results.

## 2 THEORY AND DATA ANALYSIS

For radiative shocks with negligible mixing all the energy injected into the flow as it crosses a shock front eventually appears as radiated photons as the plasma cools from its initial high temperature back down to the ambient wind temperature. While the plasma remains hotter than about  $10^6$  K most of those photons will be X-rays, whereas at lower temperatures, most of the radiation will be in the EUV, FUV, and UV, and will therefore not be observable by *Chandra* or other X-ray telescopes. For coronal plasmas the emission strength of a given line,  $\ell$ , can be characterized by a temperature-dependent emissivity,  $\Lambda_\ell(T)$ . The form of the temperature dependence arises from the dependence of the ionization balance on temperature and, to a lesser extent, from the excitation rate's temperature dependence. In Fig. 1 we show the emissivities of the lines and line complexes we analyze in the O star X-ray spectra discussed in this paper. These lines span more than an order of magnitude in temperature, but with a fair amount of overlap.

There are many additional weaker lines that contribute to *Chandra* spectra along with continuum processes – bremsstrahlung and recombination – which are relatively weak for plasmas with temperatures below 10 to 20 million K. In Fig. 2 we show the total line-plus-continuum emissivity for a coronal plasma,  $\Lambda(T)$ , according to ATOMDB (Foster et al. 2012), which is the same source we use for the individual line emissivities. We note that the ATOMDB models assume solar abundances (Anders & Grevesse 1989). More recent re-evaluations of the solar abundance (e.g. Asplund et al. 2009), as well as abundance variations among the specific program stars, would lead to factors generally of order tens of per cent adjustments to the line emissivities. We do not account for possible differences in assumed solar abundances or specific star's particular abundances, except in a few cases, for



**Figure 1.** The line emissivities from ATOMDB (Foster et al. 2012) for all the lines measured in the program stars. The four lines from each of the helium-like complexes are combined together into a single emissivity for each complex, and likewise the two components of the Ly $\alpha$  features are combined together into a single emissivity function. Several important lines at the extremes of the temperature distribution are labeled and the Fe XVII lines are denoted by the dashed, black curves. The N VI line complex is present in *XMM-Newton* spectra, but not *Chandra* spectra. We include it in the analysis of  $\zeta$  Pup only. Similarly, the dotted purple curve is for the S XVI Ly $\alpha$  line that we include only in the  $\zeta$  Pup analysis (it is not detected, but provides an interesting upper limit, as discussed in sec. 4).

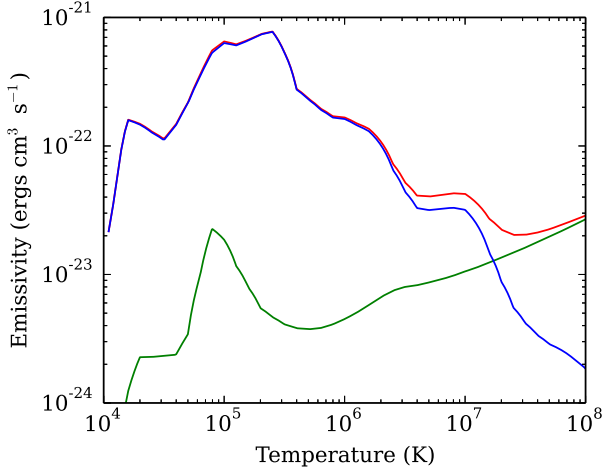
which non-solar O and N abundances are quite significant. In those cases, we simply scale the ATOMDB emissivities for the relevant lines according to the specific element's abundance. Note that traditionally in X-ray astronomy the quantity referred to as emissivity,  $\Lambda$ , has units of  $\text{ergs cm}^3 \text{s}^{-1}$ , so that multiplying it by a number density squared gives the more usual power per unit volume. The advantage of defining  $\Lambda$  in this way is that it is independent of the plasma density.

With these emissivities in hand, we can compute the fraction of the total radiated power that emerges from the plasma in a given line, at a given temperature. As the hot, post-shock plasma cools back down to its low, pre-shock value, the fraction of radiation emitted in a given line  $\ell$  for a single shock with post-shock temperature  $T_s$  is set by

$$f_\ell(T_s) = \int_0^{T_s} \frac{\Lambda_\ell(T)}{\Lambda(T)} \frac{dT}{T_s}, \quad (1)$$

where  $\Lambda_\ell(T)$  is the line emissivity (the curves shown in Fig. 1) and  $\Lambda(T)$  is the total emissivity (the upper (red) curve shown in Fig. 2)<sup>1</sup>. Note that it does not matter how quickly the plasma cools, as it will emit the same amount of energy in cooling through some temperature interval whether it does so slowly or rapidly. It is implicit in our procedure that the stars being analyzed have a large, statistical sample of shocks representing all stages of shock evolution and

<sup>1</sup> This approach is similar to the one described in §2.3 of Kee et al. (2013) for computing the fractional power radiated by a hot plasma into a particular X-ray bandpass.



**Figure 2.** The contribution of all emission lines (blue) to the total radiated power (red), along with the contribution of continuum processes (green).

therefore that a single X-ray observation is equivalent to completely tracking the evolution of a representative ensemble of shocks.

Now we consider the line emission from a distribution of shocks. Let us suppose a typical fluid parcel undergoes  $\bar{N}$  shocks in advecting out through the wind, with each shock having a cumulative probability  $p(T_s)$  for a post-shock temperature at or above  $T_s$ , which declines from unit normalization for very weak shocks,  $p(0) = 1$ . Then for a wind with mass loss rate  $\dot{M}$ , the total luminosity in the line is given by an integral over the *differential* distribution in shock temperature,  $-\frac{dp}{dT_s}$ , multiplied by the energy associated with the shocked wind mass at that temperature and the fraction of the energy that is radiated in the line (given by eqn. (1)),

$$L_\ell = \frac{5}{2} \frac{\bar{N}\dot{M}k}{\mu m_p} \int_0^\infty -\frac{dp(T_s)}{dT_s} T_s f_\ell(T_s) dT_s. \quad (2)$$

Here the mean molecular weight  $\mu$  is in units of the proton mass  $m_p$  and the post-shock enthalpy per unit mass<sup>2</sup> is  $(5/2)kT_s/\mu m_p$ . Integrating by parts and using eqn. (1) we have

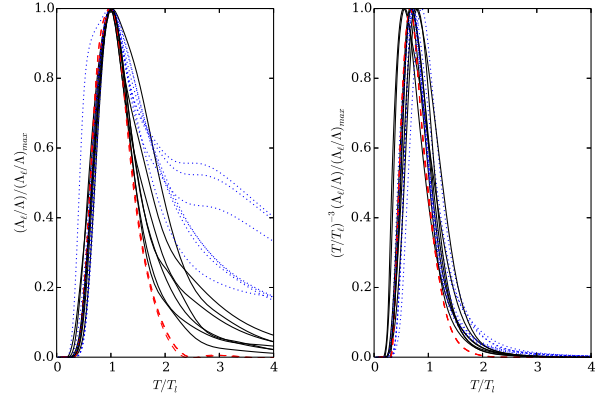
$$L_\ell = \frac{5}{2} \frac{\bar{N}\dot{M}k}{\mu m_p} \int_0^\infty \frac{\Lambda_\ell(T_s)}{\Lambda(T_s)} p(T_s) dT_s. \quad (3)$$

This equation shows that the observed line luminosity,  $L_\ell$ , depends on the convolution of the actual shock distribution function,  $p(T_s)$ , with the emissivity ratio. For notational and conceptual clarity, let us divide this convolution integral into two parts,

$$\tilde{p}_\ell \Delta T_\ell = \int_0^\infty \frac{\Lambda_\ell(T_s)}{\Lambda(T_s)} p(T_s) dT_s, \quad (4)$$

where  $\tilde{p}_\ell$ , the normalized convolution of the cumulative

<sup>2</sup> By using the 5/2 enthalpy factor rather than the 3/2 appropriate for internal energy we are allowing for the possibility that the  $p dV$  work done on the gas as it crosses the shock front could also contribute to the eventual X-ray emission.



**Figure 3.** The individual line emissivity ratios,  $\Lambda_\ell/\Lambda$ , plotted on a common temperature scale based on each line’s  $T_\ell$ , normalized to have the same peak emissivity (left) and also multiplied by a power-law of index  $n = 3$  that we find, in the next section, is representative of our program stars’ shock heating rates,  $p(T_s)$  (right). The quantity in this right-hand panel represents the key integrand in equations (3) and (4). The red, dashed curves represent iron lines, the solid, black curves, He-like ions, and the dotted, blue curves, H-like ions, with their stronger high-energy tails. Note that the peaks are not at  $T = T_\ell$  in the right-hand panel because the multiplicative power-law factor affects the emissivity ratios differently depending on their shapes.

shock distribution, is the observationally derived quantity.<sup>3</sup> The normalization factor,  $\Delta T_\ell$ , is set by atomic physics,

$$\Delta T_\ell \equiv \int_0^\infty \frac{\Lambda_\ell(T)}{\Lambda(T)} dT, \quad (5)$$

which represents the portion of the temperature change in the cooling layer that is associated with a specific line. In an analogy with spectral line equivalent width, this can be thought of as a “temperature equivalent width,” in that it is the width a line emissivity curve would have if it were rectangular in shape and accounted for all of the radiated power, rather than just a fraction, over that temperature range (Gayley 2014).

Within the assumed model, eqn. (3) is an exact integral expression for the line luminosity given emissivity functions of any form. As Fig. 1 shows, the emissivity for each line has a distinct peak at a specific temperature, and to the extent that these emissivities are strongly peaked functions of temperature, the extraction of the shock heating rate from an ensemble of line luminosity measurements is quite straightforward, as we will presently show. In Fig. 3 the individual line emissivity functions are plotted as a ratio of the total emissivity, normalized to have the same peak value, and shown on a temperature scale based on each function’s maximum (defined as  $T = T_\ell$ ). We also show these functions multiplied by a power-law of index  $n = 3$  to approximate the integrand in eqn. (4), which is the key quantity related to the observationally derived shock heating rate for each line,  $\tilde{p}_\ell$ . Given that this key functional form is quite narrow, and

<sup>3</sup> The quantity  $\tilde{p}_\ell$  is equivalent to  $\langle p \rangle_i$  defined in eqn. (18) of Gayley (2014).

for conceptual and mathematical simplicity<sup>4</sup>, here we treat the emissivity ratio as a  $\delta$ -function,

$$\frac{\Lambda_\ell(T_s)}{\Lambda(T_s)} = \Delta T_\ell \delta(T_s - T_\ell), \quad (6)$$

from which we trivially find that  $\tilde{p}_\ell = p(T_\ell)$ . Eqn. (3) then becomes

$$L_\ell = \dot{M} \frac{5k\Delta T_\ell}{2\mu m_p} \bar{N} p(T_\ell), \quad (7)$$

from which can be obtained an empirically inferred  $\bar{N} p(T_\ell)$  in terms of an observed set of line luminosities,  $L_\ell$ , and the tabulated emissivity functions that contribute to  $\Delta T_\ell$ ,

$$\bar{N} p(T_\ell) = \frac{2\mu m_p L_\ell}{5\dot{M} k \Delta T_\ell}. \quad (8)$$

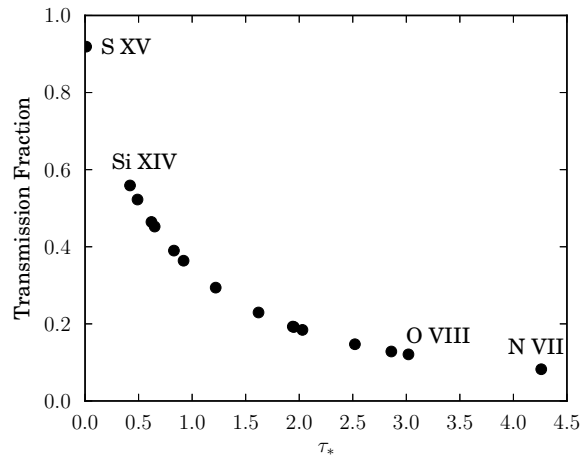
The product  $\bar{N} p(T_\ell)$  is a unitless number describing the expectation value of the number of shocks with temperature  $T_s \geq T_\ell$  that a mass parcel traverses as it flows through the wind.

The X-ray spectral data we use to make the determination of  $L_\ell$  for use in eqn. (8) are the line fluxes measured with the *Chandra* High Energy Transmission Grating Spectrometer (HETGS) (Canizares et al. 2005). We supplement these with the N VI line complex measured with *XMM-Newton* in  $\zeta$  Pup (Leutenegger et al. 2007) to provide information on the low-temperature end of the  $p(T_s)$  distribution. To convert the measured line fluxes into luminosities,  $L_\ell$ , we apply corrections for: (1) the inverse square law via the distances to the program stars; (2) the transmission of the interstellar medium; and (3) the transmission of the stellar winds themselves.

The bulk, cool component of massive-star winds is a source of continuum bound-free opacity to the EWS X-rays. Not only does this opacity lead to attenuation of the X-rays – which we must correct for – but it also leads to a characteristic asymmetry of the X-ray line profiles. The profile fitting of the observed X-ray emission lines that we use to find the line fluxes also provides information about the wind optical depth at the wavelength of each line – via the parameter,  $\tau_* \equiv \kappa \dot{M} / 4\pi R_* v_\infty$  (Owocki & Cohen 2001). Here  $\kappa$  is the wavelength-dependent bound-free opacity at the wavelength of the emission line, which is assumed to be independent of location in the wind,  $\dot{M}$  is the wind mass-loss rate,  $R_*$  is the stellar radius, and  $v_\infty$  is the stellar wind terminal velocity.

From the fitted  $\tau_*$  values, we can compute the transmission (defined as the fraction of the emitted X-ray radiation that escapes the wind) using the formalism of Leutenegger et al. (2010). We note that this transmission value is not the usual exponential form due to a slab of absorbing material between the observer and the source, but rather is a more complicated function that accounts for the spatial distribution of the emitting plasma embedded within the absorbing wind. This wind transmission correction can be significant.

<sup>4</sup> In Appendix A we explore a “power-law Boltzmann” model that accounts more realistically for the broader and asymmetric form of the emissivities. A key result is that there is still a direct connection between the observed  $\tilde{p}_\ell$  for each line and the shock distribution evaluated at the peak temperature,  $p(T_\ell)$ , differing only by a normalization factor of order unity from the  $\delta$ -function based analysis presented in this section.



**Figure 4.** The fraction of the emitted line photons that are transmitted through the wind without being absorbed, for each line in the *Chandra* spectrum of  $\zeta$  Pup, as a function of each line’s characteristic optical depth value,  $\tau_*$ , derived from fitting the line profile shapes (Cohen et al. 2014).

Fig. 4 shows the transmission values,  $T_w$ , for each of the 16 lines measured in the *Chandra* spectrum of  $\zeta$  Pup, which is the star in our sample with the most wind attenuation. Appropriate corrections are applied to the other four stars, as well, with line optical depth values taken from Cohen et al. (2014).

Given these considerations, the line luminosity,  $L_\ell$  is computed from the observed line flux,  $F_\ell$  by

$$L_\ell = 4\pi d^2 F_\ell e^{\tau_{ism}} / T_w(\tau_*), \quad (9)$$

where  $d$  is the distance to the star,  $\tau_{ism}$  is the optical depth of the interstellar medium, which we compute from the observed ISM column densities and the TBABS ISM absorption model (Wilms et al. 2000), and  $T_w(\tau_*)$  is the wind transmission. This is denoted as  $T(\tau_*)$  in Leutenegger et al. (2010), but is slightly relabeled here to better distinguish it from the temperature.

Our sample consists of the five massive stars with high quality *Chandra* grating spectra in Cohen et al. (2014) that are not contaminated by colliding wind X-ray emission and that have winds dense enough to be fully radiative out to large radii (which eliminated the weak-wind star  $\zeta$  Oph). We also exclude the O2If\* star HD 93129A which has only a small number of lines visible in its *Chandra* spectrum. For each of the five stars analyzed, Table 1 lists the relevant stellar and wind parameters. This includes the adiabatic radius,  $R_a$ , at which the radiative cooling length of a 1 keV shock equals the stellar radius (Owocki et al. 2013). Below  $R_a$ , shocks cool primarily by radiation, while above it, adiabatic expansion dominates the cooling. The bulk of the X-ray emission from embedded wind shocks in massive stars comes from the first several stellar radii of the wind according to both theoretical calculations (Feldmeier et al. 1997) and observational constraints from line profiles and forbidden-to-intercombination line intensity ratios measured via X-ray spectroscopy (Leutenegger et al. 2006). Thus the

**Table 1.** Stellar and Wind Properties

Star	Spectral Type	d (pc)	$R_*$ ( $R_\odot$ )	$v_\infty$ (km s $^{-1}$ )	$\dot{M}$ ( $M_\odot$ yr $^{-1}$ )	$R_a$ ( $R_*$ )	$N_{\text{ISM}}$ ( $10^{22}$ cm $^{-2}$ )
9 Sgr	O4 V	1300 <sup>a</sup>	12.4 <sup>b</sup>	3100	$3.7^{+1.0}_{-0.9} \times 10^{-7}$	24	0.22
$\zeta$ Pup	O4 If	460 <sup>c</sup>	18.9 <sup>d</sup>	2250	$1.76^{+0.13}_{-0.12} \times 10^{-6}$	103	0.01
$\xi$ Per	O7.5 III	382 <sup>e</sup>	14.0 <sup>f</sup>	2450	$2.2^{+0.6}_{-0.5} \times 10^{-7}$	16	0.11
$\zeta$ Ori	O9.7 Ib	226 <sup>e</sup>	22.1 <sup>b</sup>	1850	$3.4^{+0.6}_{-0.6} \times 10^{-7}$	21	0.03
$\epsilon$ Ori	B0 Ia	363 <sup>g</sup>	32.9 <sup>g</sup>	1600	$6.5^{+1.1}_{-1.5} \times 10^{-7}$	31	0.03

References: <sup>a</sup>Tothill et al. (2008); <sup>b</sup>Martins et al. (2005); <sup>c</sup>Markova et al. (2004); <sup>d</sup>Najarro et al. (2011); <sup>e</sup>van Leeuwen (2007); <sup>f</sup>Repolust et al. (2004); <sup>g</sup>Searle et al. (2008); all terminal velocities from <sup>h</sup>Haser (1995), all mass-loss rates from Cohen et al. (2014), and all ISM column densities from Fruscione et al. (1994)

values of  $R_a$  listed in the table justify the assumption of radiatively cooled shocks in the program stars.

### 3 RESULTS

The shock heating rates  $\bar{N}p(T_\ell)$  computed using equation (8) for each emission line in each of the stars are plotted in Fig. 5 as a function of the lines' temperatures of peak emissivity,  $T_\ell$ . The temperature range for each line in that figure represents the full width at half maximum (FWHM) of the line emissivity ratio,  $\Lambda_\ell(T)/\Lambda(T)$ . For each star, different lines that probe similar temperature ranges give consistent results. And there is clearly a decreasing trend for each star, consistent with the cumulative, monotonic nature of the probability function,  $p(T_s)$ . Note that this is not something imposed by our method, but rather is a reassuring consistency check on it. We note, also, that applying the wind and ISM transmission corrections improved the consistency of the results, as did accounting for lower oxygen and higher nitrogen abundances in  $\zeta$  Pup (Bouret et al. 2012; Neely et al. 2014) and  $\xi$  Per.<sup>5</sup>

Uncertainties in the derived  $\bar{N}p(T_\ell)$  values come from several different sources. The biggest uncertainties are in the stellar and wind properties – the distances and mass-loss rates – however those errors will affect every line from a given star by the same amount and so will simply scale up or down the overall shock heating rate for a given star. The sources of error that vary from line to line – statistical error on the measured brightnesses of each line and the ISM and wind transmission corrections – generally amount to a few tens of per cent. Fig. 5 shows these latter errors, but not the mass-loss rate and distance errors, as the vertical extent of the gray boxes.

We fit a power law to each star's cumulative shock heating rate – the ensemble of  $\bar{N}p(T_\ell)$  values – shown in each panel of Fig. 5 and all together in Fig. 6. The power law has the simple form

$$\bar{N}p(T_s) = N_o \left( \frac{T}{10^6 \text{ K}} \right)^{-n}. \quad (10)$$

Table 2 lists the best-fit power law model parameters,  $N_o$

**Table 2.** Power Law Fits to  $\bar{N}p(T_\ell)$  Values

Star	Spectral Type	$N_o$	$n$
9 Sgr	O4 V	0.90	2.38
$\zeta$ Pup	O4 If	0.26	2.20
$\xi$ Per	O7.5 III	0.50	3.02
$\zeta$ Ori	O9.7 Ib	0.53	3.24
$\epsilon$ Ori	B0 Ia	0.14	2.84

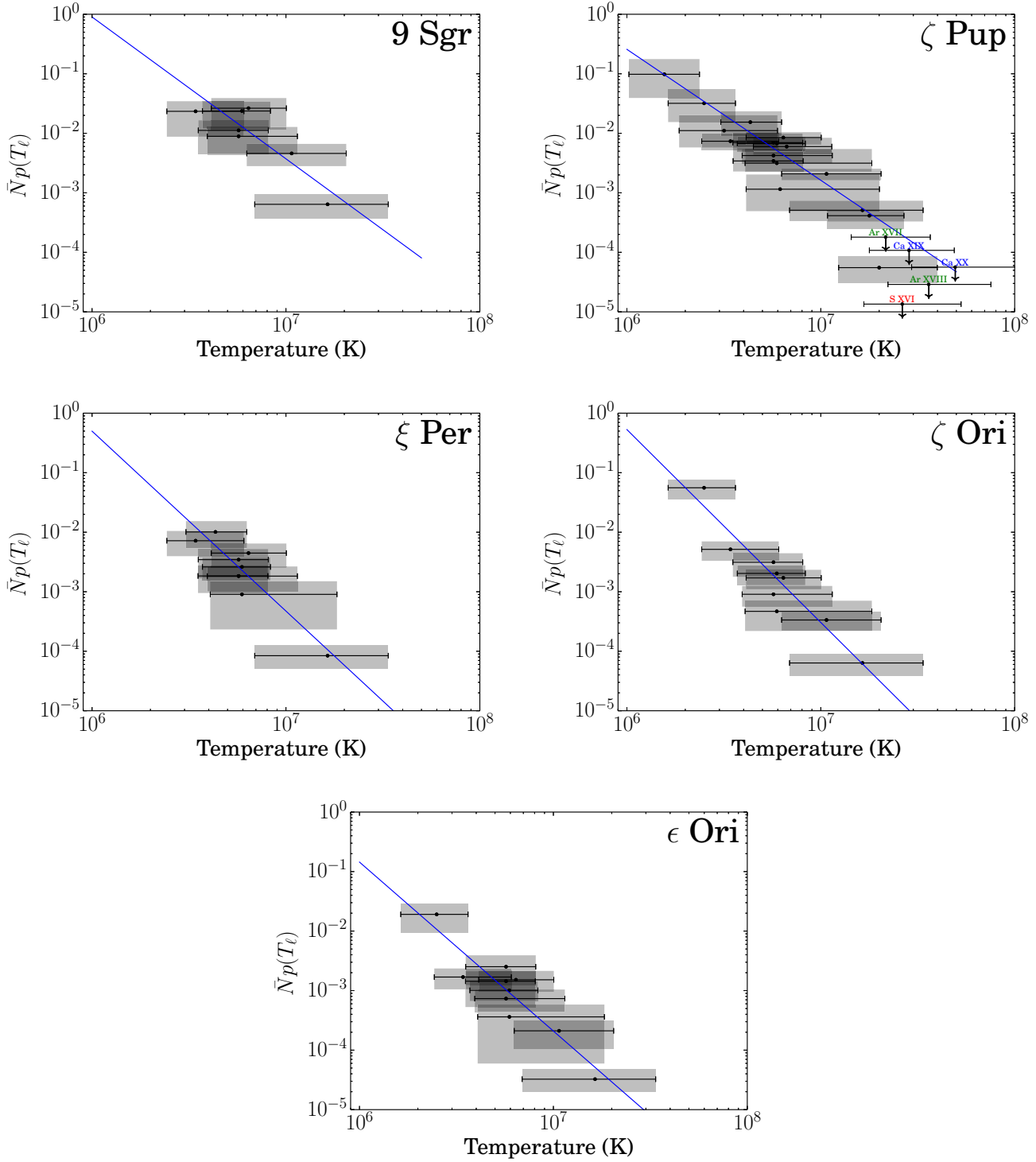
and  $n$ . All of the measured lines have peak emissivity ratios at temperatures,  $T_\ell > 10^6$  K, so determining a value for  $N_o$  requires some extrapolation. A less model-dependent statement can be made looking at Fig. 6 – all the stars have  $\bar{N}p(T)$  values approaching 0.1 for their coolest lines, which have peak temperatures,  $T_\ell$ , between 2 and  $4 \times 10^6$  K. For  $\zeta$  Pup, the N VI complex observed with *XMM-Newton* probes temperatures near  $1.5 \times 10^6$  K and has an  $\bar{N}p(T_\ell)$  value slightly above 0.1.

### 4 DISCUSSION

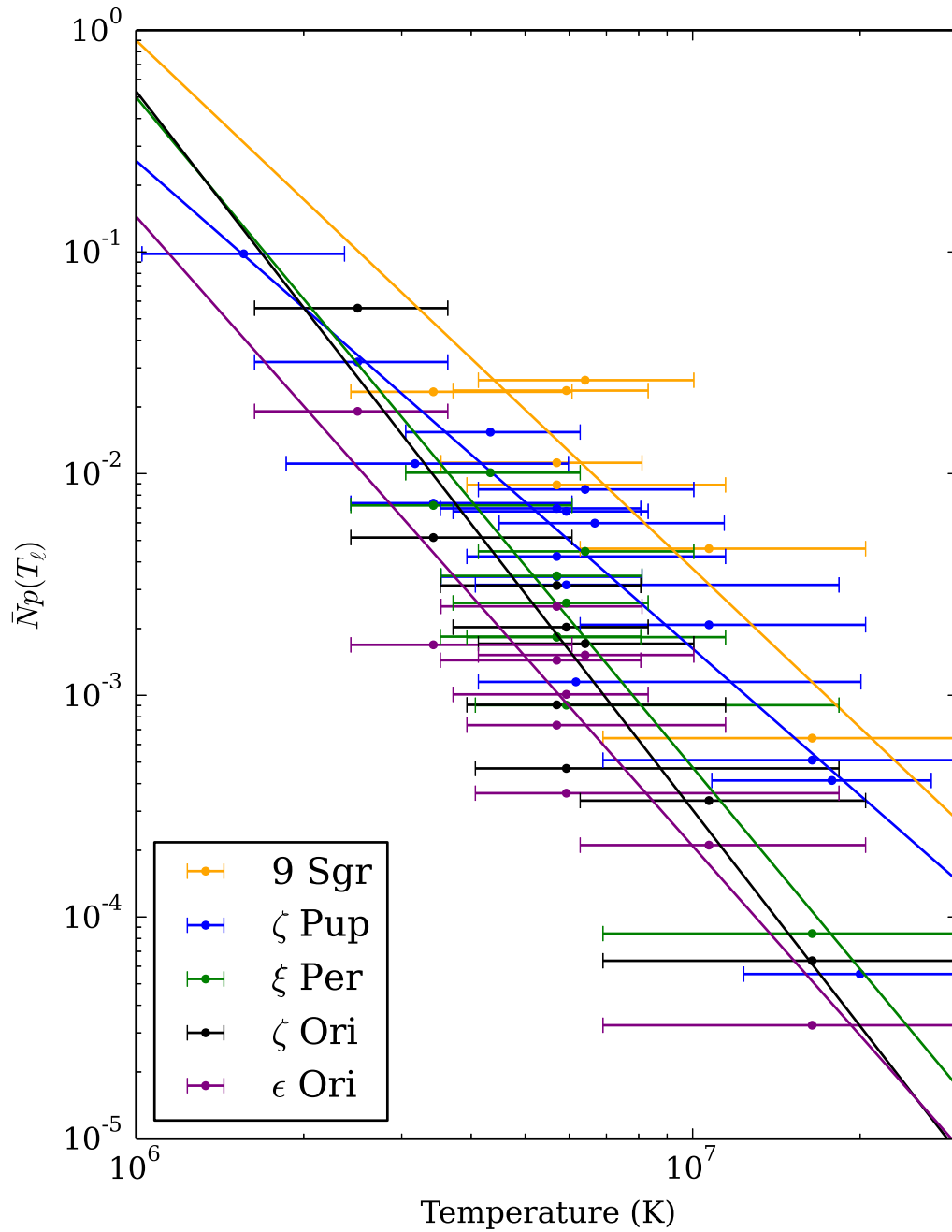
The results for all five stars are quite similar, with the  $\bar{N}p(T_\ell)$  shock heating rate consistent with a power law index of roughly  $n = 3$  and about ten per cent of the wind mass passing through a shock of  $T_s > 2 \times 10^6$  K. This temperature is roughly the minimum plasma temperature that produces significant radiation in the X-ray bandpass (ud-Doula et al. 2014), and so our values of  $N_o$  reveal that much or even most of the wind contributes to the observed X-ray emission from these stars, although simulations and observations both show that, at any given instant, only a small fraction of the wind mass is hot enough to emit X-rays. This is, however, consistent with  $N_o \approx 1$  since the typical shock cooling time is much shorter than the wind flow time.

If we consider the differential probability distribution of shock strengths, rather than the directly derived cumulative one, then we have an even steeper slope of  $T^{-4}$ , implying that a shock is  $10^4$  times less likely to heat a mass parcel of wind plasma to within 1 K of  $T = 10^7$  K than it is to heat it to within 1 K of  $T = 10^6$  K. If we consider heating per decade then we are back to having a slope of  $T^{-3}$ ; that is  $\frac{d\bar{N}p}{d \log T} \propto T^{-3}$ . So a shock is 1000 times less likely to heat a mass parcel to within a fixed fraction of  $T = 10^7$  K as it is to heat it to within that same fixed fraction of  $T = 10^6$  K.

<sup>5</sup> For this latter star, no nitrogen lines are used, and we make a factor of 1/3 correction to the oxygen abundance, which we base on visual inspection of the X-ray spectrum.



**Figure 5.** The shock heating rate,  $\bar{N}p(T_\ell)$ , is shown with the uncertainty on its value (vertical extent of each gray box) as well as the FWHM of the line emissivity ratio,  $\Lambda_\ell(T)/\Lambda(T)$  (horizontal extent, visually reinforced by the horizontal error bars). The points are at  $T_\ell$  for each line. The best-fit power law to each set of values is shown as a blue line in each panel. For  $\zeta$  Pup, the lowest temperature point corresponds to the N VI feature measured with *XMM-Newton*. And for this star, upper limits are included for five additional lines, none of which are detected in the *Chandra* spectrum.



**Figure 6.** The shock heating rate,  $\bar{N}p(T_\ell)$ , derived from each line of each of our program stars. These are the same results shown in Fig. 5 but simply collected together to facilitate comparison. We do not show the uncertainties on each point – corresponding to the vertical extent of the gray boxes – to keep the plot from being too cluttered. Because the lower range of the x-axis is  $T = 10^6$  K, the power-laws’ y-intercepts give the value of  $N_o$  for each star, as defined in eqn. (10).



The lines from the hottest plasma tend to be weak, as the strongly decreasing derived shock distributions indicate. We have little direct information about plasma with temperature much in excess of  $10^7$  K and, specifically, it is difficult to know whether there is a high-temperature cut-off to the plasma temperature distribution. Bear in mind that the power-laws we show in Fig. 6 are fits to the data-derived shock heating rates for each line, but we have not attempted to show that the distributions truly are power laws. Consulting Fig. 1, we can see that the Si XIV line complex near  $6 \text{ \AA}$  does not have significant emissivity much below  $10^7$  K, though that line is present in the high signal-to-noise spectrum of  $\zeta$  Pup only, while the Mg XII Ly $\alpha$  line with a slightly lower temperature response is present in each of our program stars. So, it does seem safe to say that the shock heating distribution reliably extends to at least  $10^7$  K.

To further explore the constraints on the hottest plasma and the strongest shocks, for the star  $\zeta$  Pup we have extracted upper limits for five high-temperature lines that should be at least moderately strong in coronal plasmas with temperature in excess of  $10^7$  K. We then compute upper limits on the quantity  $\bar{N}p(T_\ell)$  given the measured upper limits for the line luminosities,  $L_\ell$ . These five limits are shown in the  $\zeta$  Pup panel in Fig. 5 as downward facing arrows. The lowest of these, from the S XVI Ly $\alpha$  line, probing temperatures between about 15 and 50 million K, is somewhat interesting, lying about an order of magnitude below the fitted power law shock heating rate,  $\bar{N}p(T_s)$ . Similarly, the shock heating rate derived from the measured Si XIV line, lies substantially below the power law fit. As we show in the Appendix, for hydrogen-like ions such as these, the empirical shock heating rate derived for a particular line using eqn. (4),  $\tilde{p}_\ell$ , may underestimate the true shock heating rate for the line,  $p(T_\ell)$ , by about a factor of two. Correcting for this effect would still leave those two points substantially below the power law fit, suggesting – but not providing conclusive evidence – that there is a strong decline, or even a cut-off, in the shock heating rate at temperatures above about  $10^7$  K.

The results derived here for the shock heating rates can be compared to the results from the more traditional DEM approach (Wojdowski & Schulz 2005) which finds  $\frac{dEM}{dT} \sim T^{-2}$  for normal O stars generally (and two of our program stars,  $\zeta$  Pup and  $\zeta$  Ori, specifically). The emission measures reported on a line-by-line basis in Wojdowski & Schulz (2005) show significantly more scatter than our  $\bar{N}p(T_\ell)$  results do. This is likely due partially to those author’s neglect of wind absorption, which is difficult to account for in the traditional DEM approach. And in the case of  $\zeta$  Pup it is also likely due to the neglect of non-solar abundances in the DEM analysis. Finally the steeper slope we find for  $\bar{N}p(T_s)$  can be reconciled with the shallower overall trend Wojdowski & Schulz (2005) report for the O star DEMs because the differential emission measure should approximate the ratio of the heating rate to the cooling rate, and since radiative cooling in the  $10^6 < T < 10^7$  K range is a modestly decreasing function of temperature, the heating rate should indeed have a steeper (negative) slope than the emission measure.

Finally, we emphasize that our shock heating rate approach provides a physically meaningful overall normaliza-

tion, namely the expectation value for the number of shocks a typical wind mass parcel passes through, whereas the DEM provides only a snapshot of the amount of wind material emitting X-rays at any given time. Interpreting that quantity in terms of the physically more informative shock heating rate would require modeling the cooling as well as the heating, which depends on the assumptions that are made about the local density in the post-shock X-ray emitting volumes for this density-squared diagnostic.

Our new results present specific targets for simulations of embedded wind shocks. They imply an efficient shock heating mechanism, but one that is a strongly decreasing function of shock temperature and which rarely produces shocks hotter than  $10^7$  K. And they also strongly suggest that the shock heating mechanism’s characteristics are not too sensitive to stellar or wind parameters, as the results for our program stars are relatively uniform, despite a wide range of spectral subtypes, wind mass-loss rates, and terminal velocities. To the extent that there are differences among the program stars’ results, there appears to be a higher shock-heating rate for the stars with earlier spectral subtypes. Fig. 6 shows that in the middle of the temperature range sampled by the lines we observe, there is roughly an order of magnitude range in the overall shock heating rate levels among the sample stars.

It will be interesting to see if the self-excited LDI can reproduce the observational results we have derived here or if perturbations at the wind base and the associated clump-clump collisions (Feldmeier et al. 1997; Sundqvist & Owocki 2013) will be required to explain the results. Although there are few relevant predictions in the current literature, numerical simulations of the self-excited instability show velocity dispersions of a few  $100 \text{ km s}^{-1}$  (Runacres & Owocki 2002; Dessart & Owocki 2005), representing how much variation the wind velocity shows at a given radius over a long simulation run-time. This velocity dispersion is largely due to shocks but its magnitude depends on the duty cycle of shocks as well as their strengths, and so it is difficult to interpret directly in terms of a shock heating rate. The typical velocity dispersion value corresponds to a shock temperature of roughly  $10^6$  K ( $T_{\text{shock}} \approx 10^6 (v_{\text{shock}}/300 \text{ km s}^{-1})^2 \text{ K}$ ), consistent with the dominance of weak shocks over strong shocks that we find in this paper. In principle, for one-dimensional numerical simulations especially, it should be possible to track shock fronts and empirically determine both the mass flux and the characteristic shock temperature for each shock and in that way compute the  $\bar{N}p(T_s)$  predicted by theory under various assumptions. We plan to examine this in future work.

## ACKNOWLEDGMENTS

Support for this work was provided by the National Aeronautics and Space Administration through ADAP grant NNX11AD26G and *Chandra* grants AR2-13001A and TM3-14001B to Swarthmore College, ATP grant NNX11AF83G to University of Iowa, and ATP grant NNX11AC40G to University of Delaware. JOS was supported by DFG grant Pu117/8-1. ZL was supported by the Physics and Astronomy Department and the Provost’s Office of Swarthmore College via a Vanderveld-Cheung Summer Research Fellowship.

We thank Randall Smith for his help and advice regarding ATOMDB, and we thank Marc Gagné for useful discussions about 9 Sgr.

## REFERENCES

- Anders E., Grevesse N., 1989, *Geo. Cosm. Acta*, 53, 197  
 Antokhin I. I., Owocki S. P., Brown J. C., 2004, *A&A*, 611, 434  
 Asplund M., Grevesse N., Sauval A. J., Scott P., 2009, *ARAA*, 47, 481  
 Bouret J. C., Hillier D. J., Lanz T., Fullerton A. W., 2012, *A&A*, 544, 67  
 Canizares C. R., et al., 2005, *PASP*, 117, 1144  
 Cassinelli J. P., Miller N. A., Waldron W. L., MacFarlane J. J., Cohen D. H., 2001, *ApJ*, 554, L55  
 Cohen D. H., Wollman E. E., Leutenegger M. A., Sundqvist J. O., Fullerton A. W. A., Zsargó J., Owocki S. P., 2014, *MNRAS*, 439, 908  
 Cooper R. G., 1994, University of Delaware  
 Dessart L., Owocki S. P., 2003, *A&A*, 406, L1  
 Dessart L., Owocki S. P., 2005, *A&A*, 437, 657  
 Feldmeier A., 1995, *A&A*, 299, 523  
 Feldmeier A., Puls J., Pauldrach A. W. A., 1997, *A&A*, 322, 878  
 Foster A. R. Ji L., Smith R. K., Brickhouse N. S., 2012, *ApJ*, 756, 128  
 Fruscione A., Hawkins I., Jelinsky P., Wiercigroch A., 1994, *ApJS*, 94, 127  
 Gayley K. G., 2014, *ApJ*, submitted  
 Haser S. M., 1995, Universitäts-Sternwarte der Ludwig-Maximilian Universität, München  
 Hervé A., Rauw G., Nazé Y., Foster A., 2012, *ApJ*, 748, 89  
 Kaastra J. S., Mewe R., Liedahl D. A., Singh K. P., White N. E., Drake S. A., 1996, *A&A*, 314, 547  
 Kahn S. M., Leutenegger M. A., Cottam J., Rauw G., Vreux J.-M., den Boggen A. J. F., Mewe R., Güdel M., 2001, *A&A*, 365, L312  
 Kee , Owocki S. P., ud-Doula A., 2013, *MNRAS*, 438, 3557  
 Kramer R. H., Cohen D. H., Owocki S. P., 2003, *ApJ*, 592, 532  
 Leutenegger M. A., Cohen D. H., Zsargó J., Martell E. M., MacArthur J. P., Owocki S. P., Gagné M., Hillier D. J., 2010, *ApJ*, 719, 1767  
 Leutenegger M. A., Owocki S. P., Kahn S. M., Paerels F. B. S., 2007, *ApJ*, 659, 642  
 Leutenegger M. A., Paerels F. B. S., Kahn S. M., Cohen D. H., 2006, *ApJ*, 650, 1096  
 Liefke C., Ness J.-U., Schmitt J. H. M. M., Maggio A., 2008, *A&A*, 491, 859  
 Lucy L. B., Solomon P. M., 1970, *ApJ*, 159, 879  
 Markova N., Puls J., Repolust T., Markov H., 2004, *A&A*, 413, 693  
 Martins F., Schaerer D., Hillier D. J., 2005, *A&A*, 436, 1049  
 Milne E. A., 1926, *MNRAS*, 86, 459  
 Najarro F., Hanson M. M., Puls J., 2011, *A&A*, 535, 32  
 Neely J. E., Cohen D. H., Leutenegger M. A., Sundqvist J. O., Petit V., Owocki S. P., 2014, in prep.  
 Owocki S. P., Castor J. I., Rybicki G. B., 1988, *A&A*, 335, 914  
 Owocki S. P., Cohen D. H., 2001, *ApJ*, 559, 1108

- Owocki S. P., Sundqvist J. O., Cohen D. H., Gayley K. G., 2013, *ApJ*, 429, 3379  
 Peterson J. R., Fabian A. C., 2006, *Phys. Rep.*, 427, 1  
 Repolust T., Puls J., Herrero A., 2004, *A&A*, 415, 349  
 Runacres M. C., Owocki S.P., 2002, *A&A*, 381, 1015  
 Searle S. C., Prinja R. K., Massa D., Ryan R., 2008, *A&A*, 481, 777  
 Sundqvist J. O., Owocki S. P., 2013, *MNRAS*, 428, 1837  
 Tothill, N. F. H., Gagné, M., Stecklum, B., Kenworthy, M. A., 2008, *Handbook of Star Forming Regions*, Volume II, 533  
 ud-Doula A., Owocki S. P., Townsend R. H. D., Petit V., Cohen D. H., 2014, *MNRAS*, in press  
 van Leeuwen F., 2007, *A&A*, 474, 653  
 Wilms J., Allen A., McCray R., 2000, *ApJ*, 542, 914  
 Wojdowski P. S., Schulz N. S., 2005, *ApJ*, 627, 935  
 Zhekov S. A., Palla F., 2007, *MNRAS*, 382, 1124

## APPENDIX A: POWER-LAW BOLTZMANN MODEL FOR LINE EMISSION

To infer the wind shock distribution from observed X-ray emission line fluxes, §2 assumes the emissivity ratio  $\Lambda_\ell(T)/\Lambda(T)$  can be modeled as a narrow,  $\delta$ -function form about a peak emission temperature  $T_\ell$ . But the plot in Fig. 1 suggests instead that this line emissivity has a sharp, exponential, Boltzmann-factor cutoff at low  $T$ , and a more gradual power-law decline at high  $T$ . Let us thus examine the effects of such a “Power-Law Boltzmann” (PLB) form for the line emissivity.

Specifically, for the case of a Boltzmann cutoff with an asymptotic power-index  $q > 0$  at high temperature, let us write

$$\frac{\Lambda_\ell(T)}{\Lambda(T)} \approx C_\ell e^{-qT_\ell/T} \left(\frac{T}{T_\ell}\right)^{-q}, \quad (\text{A1})$$

which has its peak at temperature  $T = T_\ell$ , with  $C_\ell$  a normalization constant to be set below. For a shock distribution that is itself a pure power law with index  $n$ , i.e.

$$p(T_s) = \left(\frac{T_s}{T_o}\right)^{-n}, \quad (\text{A2})$$

the convolution integral in eqn. (3) takes the form

$$\begin{aligned} \int_0^\infty \frac{\Lambda_\ell(T_s)}{\Lambda(T_s)} p(T_s) dT_s &= C_\ell T_\ell q^{1-n-q} \Gamma(n+q-1) \left(\frac{T_\ell}{T_o}\right)^{-n} \\ &\equiv C_\ell T_\ell K_{nq} p(T_\ell); \quad n+q > 1, \quad (\text{A4}) \end{aligned}$$

where  $\Gamma$  is the complete Gamma function, and the restriction  $n+q > 1$  is required to insure convergence of the integral. In defining a constant factor  $K_{nq}$  that depends on the power exponents  $n$  and  $q$ , the latter equality (A4) shows that the convolution over this more general PLB form for the emission still yields the original shock distribution at peak emission temperature,  $p(T_\ell)$ , but now just with a somewhat different normalization.

Note in particular that for  $n = 0$  and thus  $q > 1$ , the integral (A3) reduces to the definition (5), giving now  $\Delta T_\ell = C_\ell T_\ell K_{0q}$ . For cases with  $q > 1$ , we can use this to eliminate the normalization constant  $C_\ell$  in favor of  $\Delta T_\ell$ , yielding

$$\tilde{p}_\ell \equiv \int_0^\infty \frac{\Lambda_\ell(T_s)}{\Lambda(T_s)} p(T_s) \frac{dT_s}{\Delta T_\ell} \quad (\text{A5})$$

$$= \frac{\Gamma(n+q-1)}{q^n \Gamma(q-1)} p(T_\ell) \quad (\text{A6})$$

$$\equiv \frac{K_{nq}}{K_{0q}} p(T_\ell) ; \quad q > 1, \quad (\text{A7})$$

which thus provides a convenient notation for comparison with results from §2.

Rather remarkably, we thus find that, for this PLB form for line emission, the emission-weighted integration over the shock distribution,  $\tilde{p}_\ell$ , still reproduces the actual shock distribution,  $p(T_\ell)$ , with now just an added renormalization factor.

Thus, at least in the case that the original shock distribution is indeed a pure power law, all the previous results – derived under the assumption that the emission has a narrow,  $\delta$ -function form about the peak – can still be retained for this more realistic PLB form, *if* one simply makes a modest renormalization in the overall level of the inferred distribution!

As a specific example, for He-like ions, like O VII or N VI, we find  $q \approx 3$ . For a typical inferred shock distribution power-index  $n \approx 3$ , this gives

$$\tilde{p}_\ell = \frac{8}{9} p(T_\ell) ; \quad n = 3, q = 3, \quad (\text{A8})$$

showing that in this case the renormalization correction is just slightly below unity.

For H-like ions like O VIII or Si XIV, we find  $q \approx 1.7$ , giving,

$$\tilde{p}_\ell = 0.65 p(T_\ell) ; \quad n = 3, q = 1.7, \quad (\text{A9})$$

showing that in this case the renormalization correction is moderately below unity.