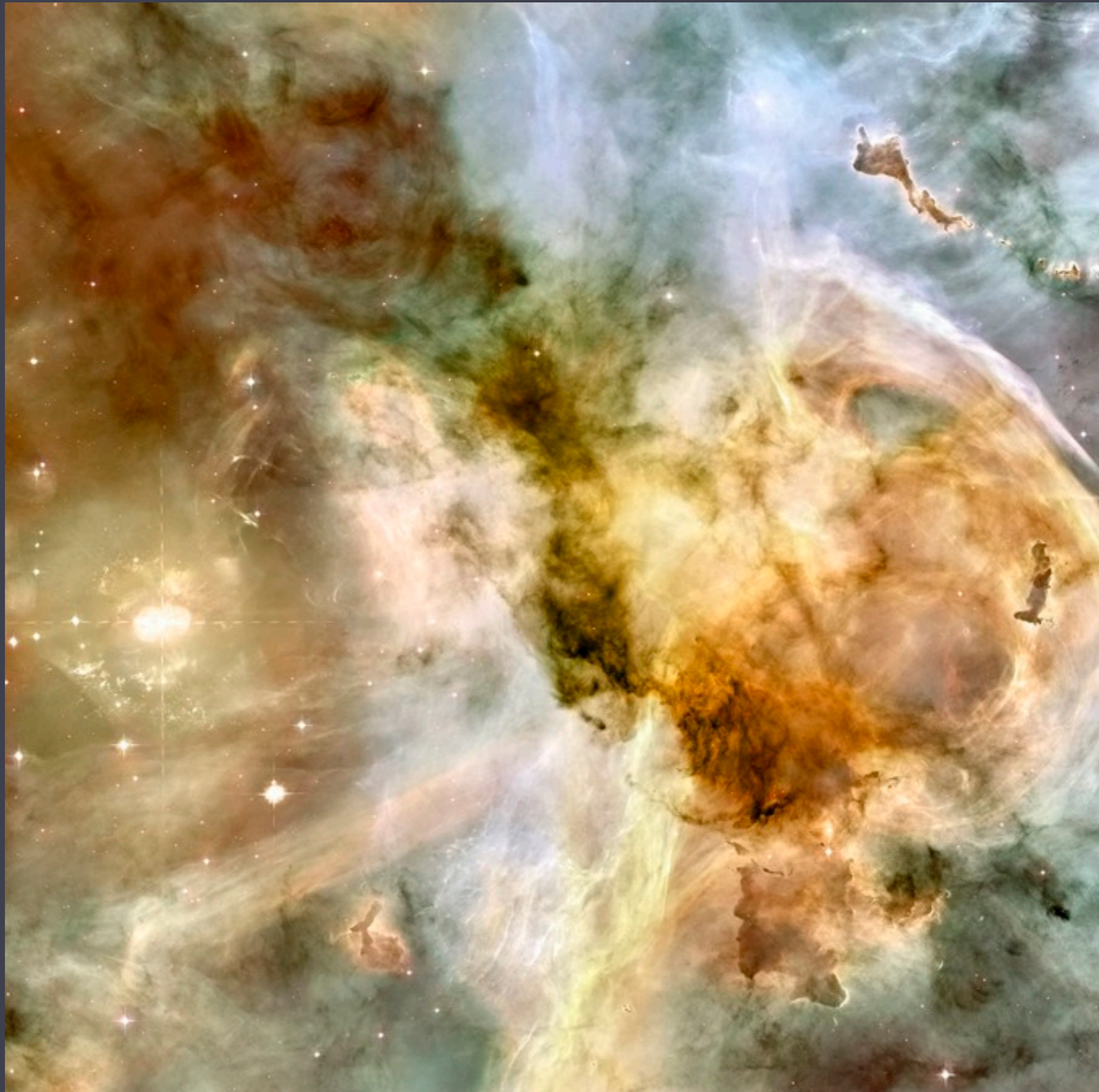


Carina Nebula: star formation region, $d \sim 7000$ light years



HST: Carina Nebula

energized by the few dozen most massive & luminous stars



HST: Carina Nebula

massive stars produce heavy elements and return them to the Galaxy via their stellar winds

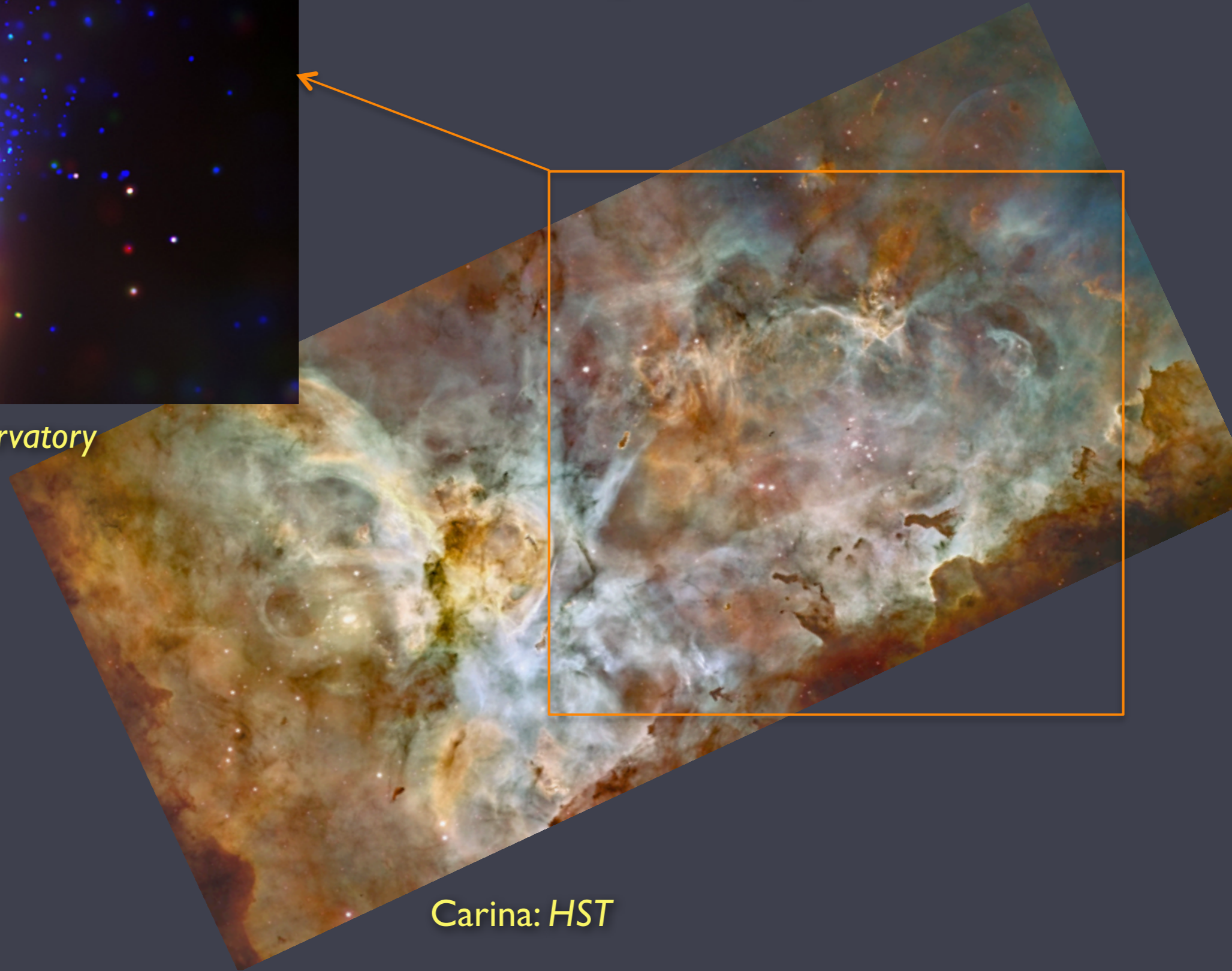


eta Carina

The massive stars are also
strong X-ray sources



Tr 14: Chandra X-ray Observatory



Carina: HST

***Chandra* X-ray spectroscopy of the very early O supergiant HD 93129A: constraints on wind shocks and the mass-loss rate**

David H. Cohen,^{1*} Marc Gagné,² Maurice A. Leutenegger,^{3,4} James P. MacArthur,¹ Emma E. Wollman,^{1,5} Jon O. Sundqvist,⁶ Alex W. Fullerton⁷ and Stanley P. Owocki⁶

¹*Department of Physics and Astronomy, Swarthmore College, Swarthmore, PA 19081, USA*

²*Department of Geology and Astronomy, West Chester University, West Chester, PA 19383, USA*

³*NASA/Goddard Space Flight Center, Code 662, Greenbelt, MD 20771, USA*

⁴*CRESST and University of Maryland, Baltimore County, MD 21250, USA*

⁵*Department of Physics, Caltech, 1200 East California Boulevard, Pasadena, CA 91125, USA*

⁶*Bartol Research Institute, University of Delaware, Newark, DE 19716, USA*

⁷*Space Telescope Science Institute, 3700 San Martin Drive, Baltimore, MD 21218, USA*



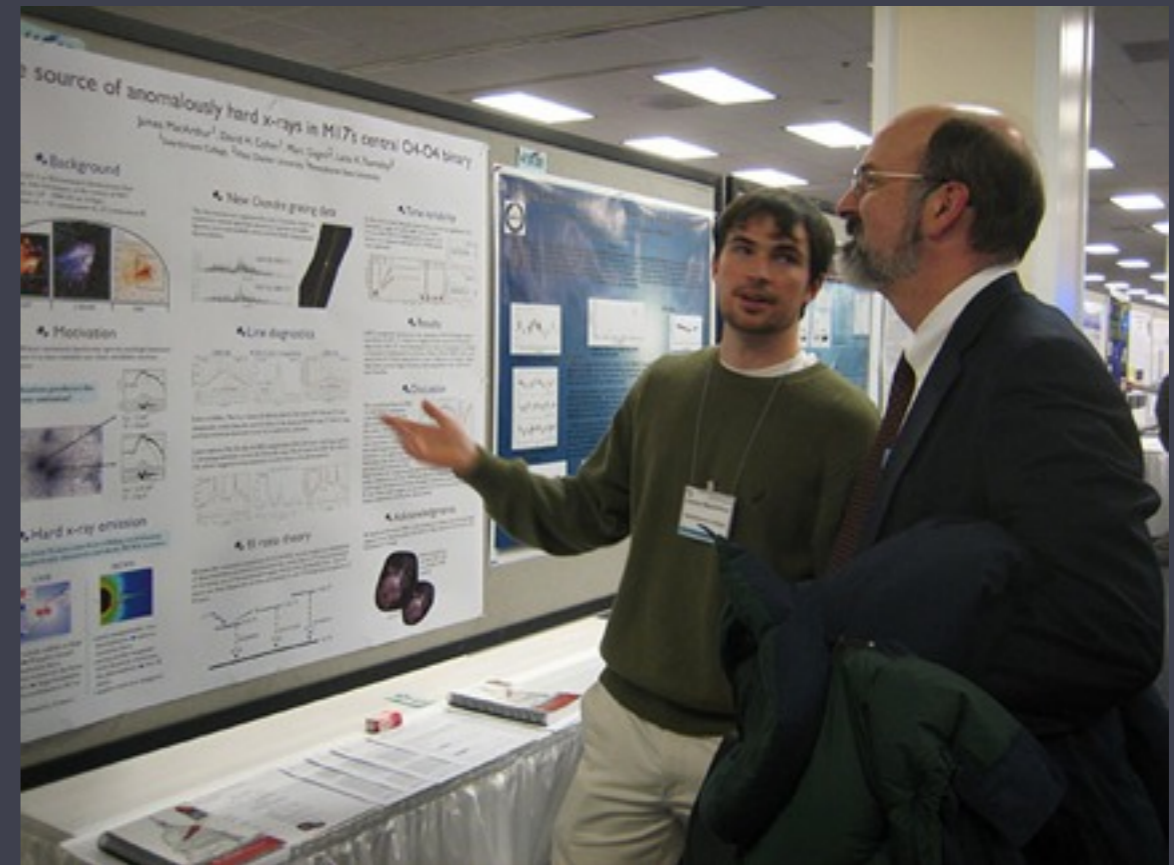
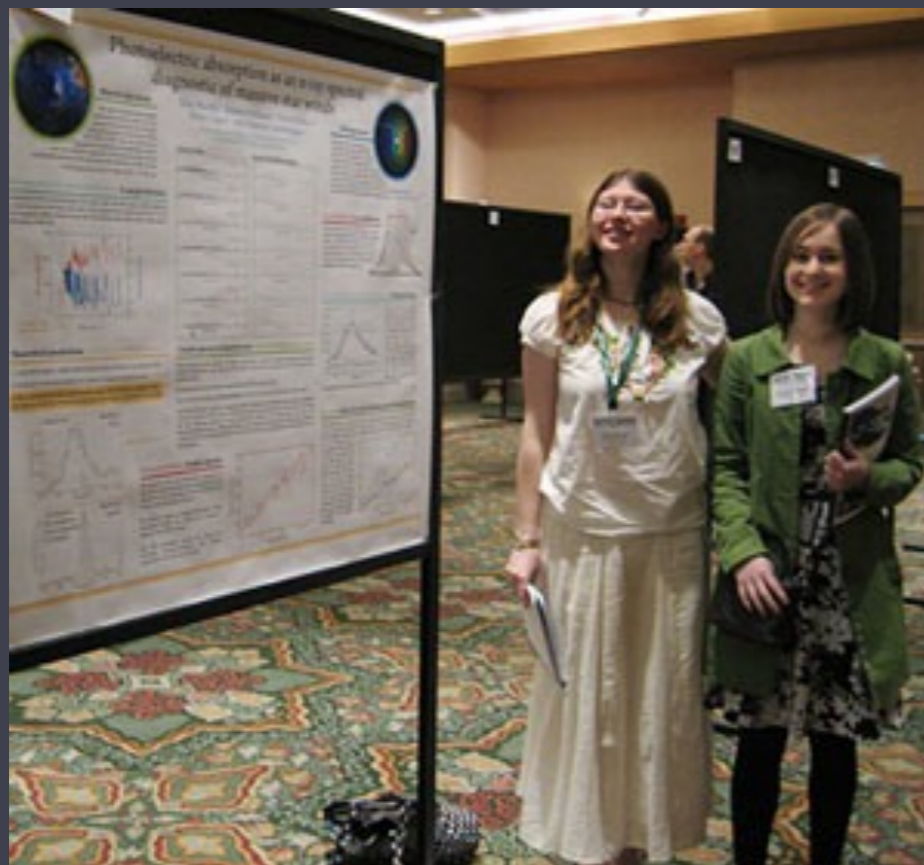
X-ray spectroscopy of the most massive star in the cluster has already yielded new information about this star's wind

Tr 14 in Carina: *Chandra* X-ray Observatory

Working this summer will require:

1. Willingness to work a lot with computers - some programming, a lot of using hard-to-use programs.
2. Enthusiasm for trying things until you find something that works.
3. But also then carefully applying the useful technique and being very organized about it.
4. Dedication to getting things right.
5. Motivation to read papers, learn about the context of the problems you're working on.

science is a
social activity



more information on my website

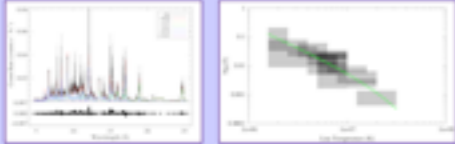
astro.swarthmore.edu/~cohen

all faculty websites are listed on the dept site


Student Research Group

Summer Research 2013

Zack Li and Kelley Langhans, both rising sophomores, are working on different aspects of the problems of measuring and modeling the plasma temperature distribution in the shock-heated winds of massive stars.



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Research Links

Astrobetter - lots of useful information about the nitty-gritty of astronomical research

- ADS
- astro-ph
- SIMBAD
- Astronomical Catalogs
- SkyView
- Chandra X-ray Center
- XMM Guest Observer Facility
- HEASARC
- ATOMDB atomic database
- physical and astronomical constants
- astrophysical constants and data

Graphics

- Historical graphics
- Edward Tufte
- Infographics News
- Visualizing Astronomy at the CFA
- Information Aesthetics

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- HHMI travel funding
- Sigma Xi travel funding (these two will fund travel to meetings)
- Sigma Xi grants in aid of research
- National Geographic Young Explorers (things like travel to an observatory are eligible)
- DoD SMART scholarships
- NDSEG graduate student scholarship
- Society of Physics Students
- American Astronomical Society

Grad School Information

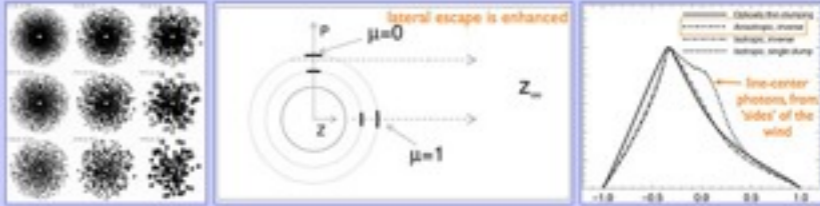
- Choosing a graduate program
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Recent Presentations

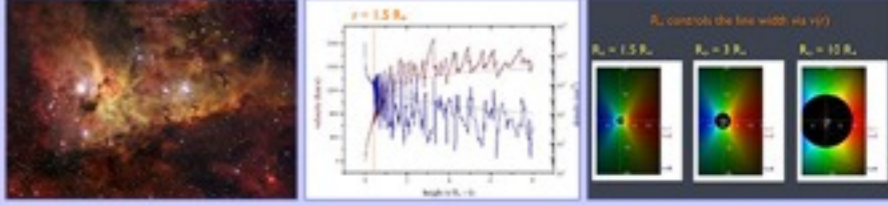
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At the meeting to honor Tony Moffat, in Lac Taureau, Quebec, I gave a talk on HD 93129A, one of the most massive stars in the Galaxy [pdf of talk and pdf of paper].



...and there are a few more slides with examples of the kinds of problems we work on (and sometimes solve) following this slide

Two outstanding questions

1. How do massive stars produce their strong X-ray emission? What's the physics of the wind shocks?
2. How strong are their stellar winds (what are their mass-loss rates)?

observational X-ray astronomy

X-ray spectroscopy with the Chandra X-ray Telescope



Chandra in the Space Shuttle cargo bay

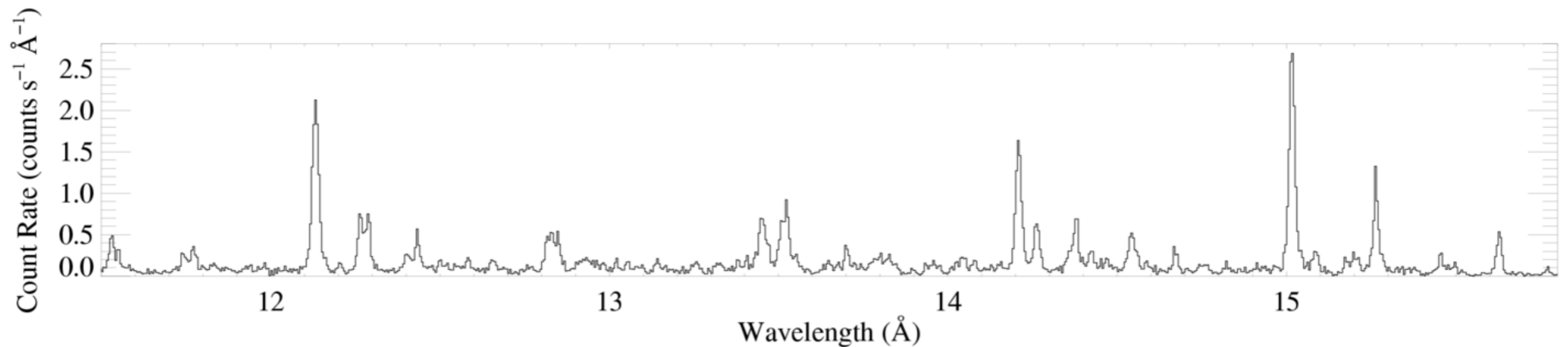
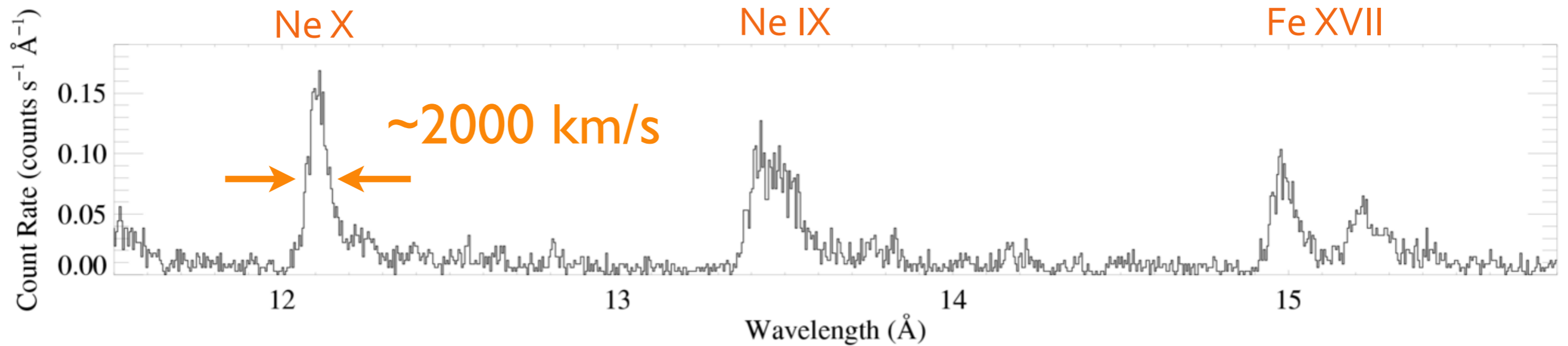
Chandra in orbit



massive stars' X-ray emission lines are **broad**

Chandra spectra

ζ Pup (O4If)



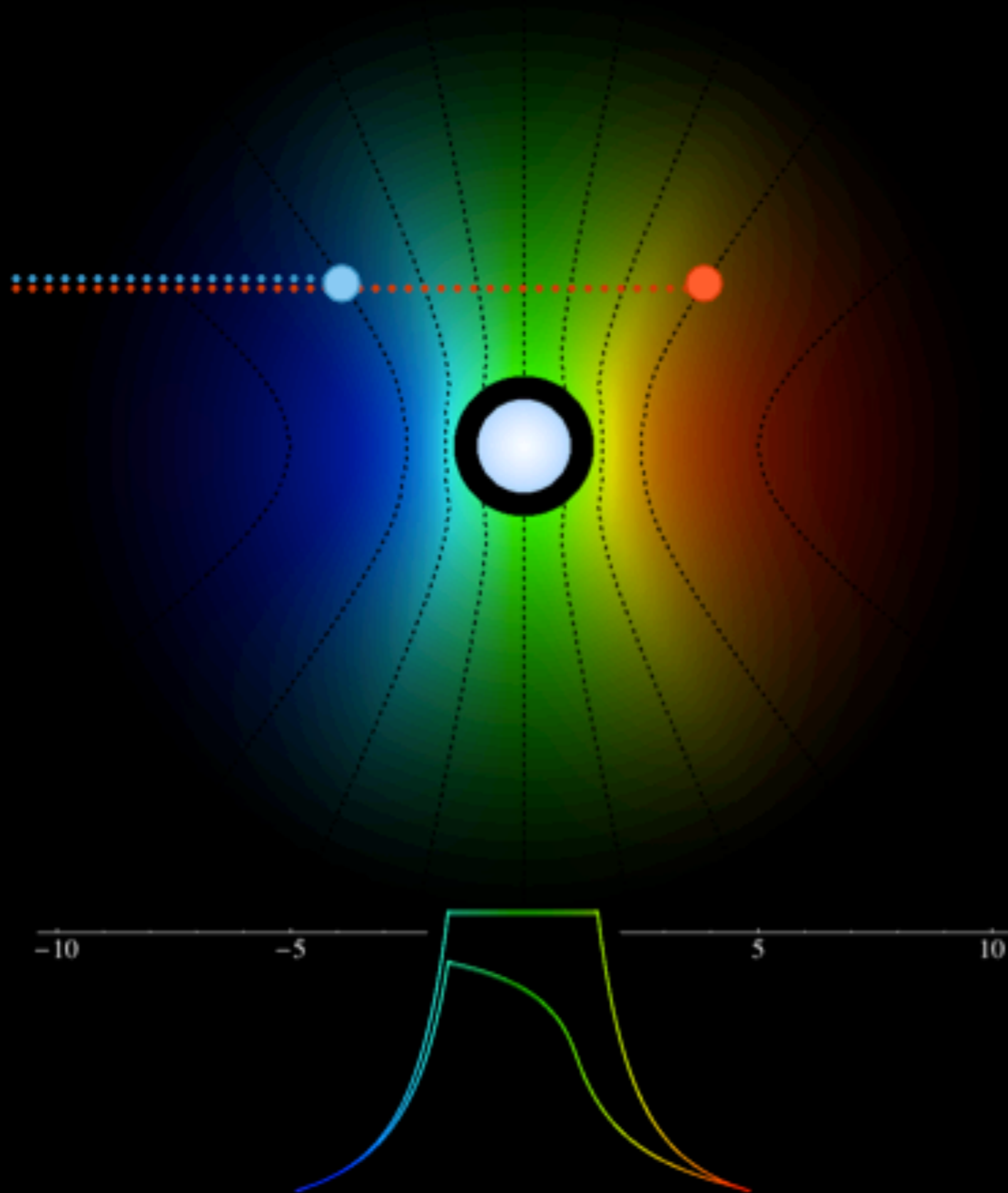
Capella: G star for comparison (narrow lines)

we make models

Line Asymmetry

$$\tau = \tau_* \int_z^\infty \frac{R_* dz'}{r'^2 (1 - R_*/r')^\beta}$$

A

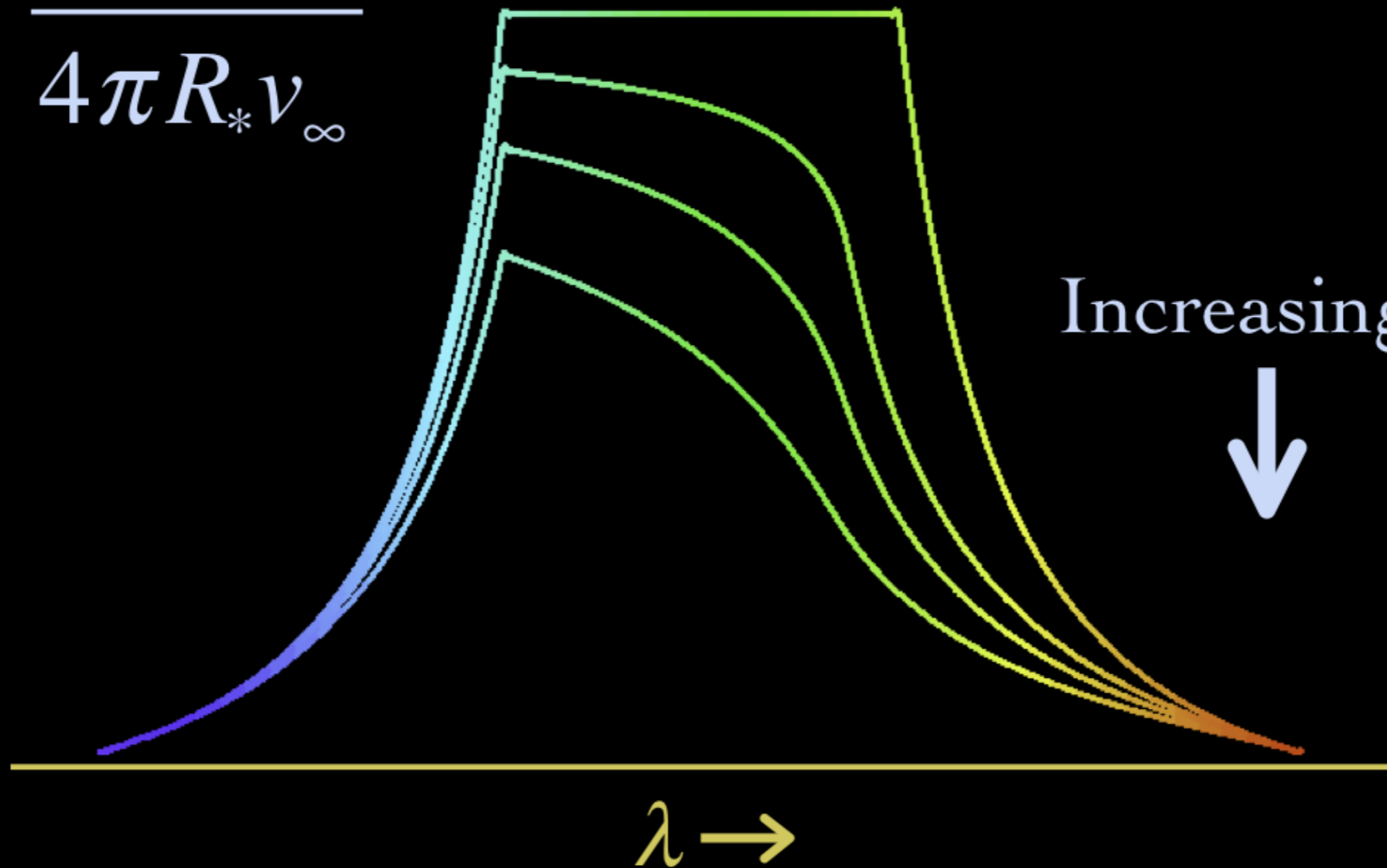


models make predictions

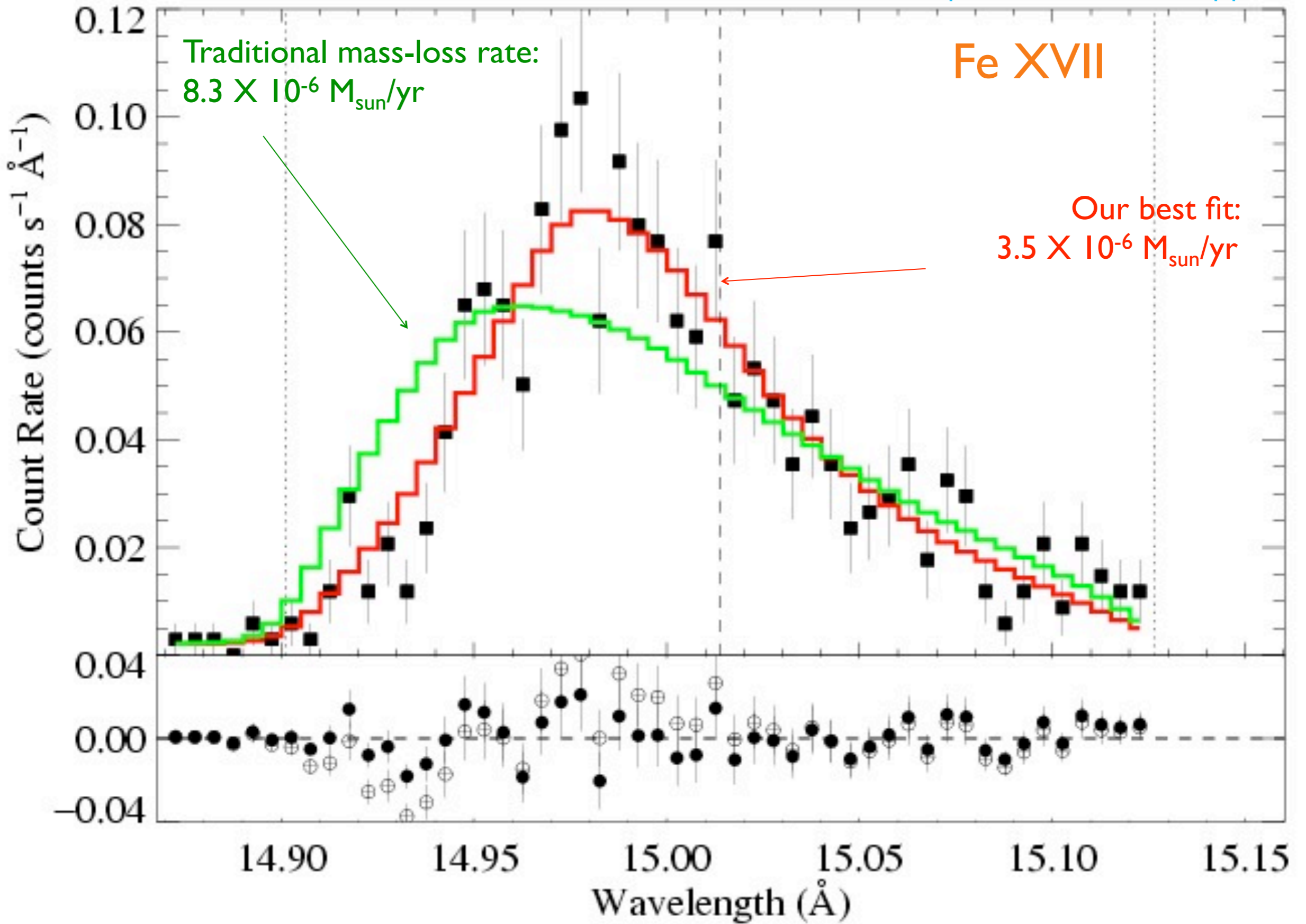
Wind Profile Model

wind mass-loss rate

$$\tau_* = \frac{\kappa \dot{M}}{4\pi R_* v_\infty}$$



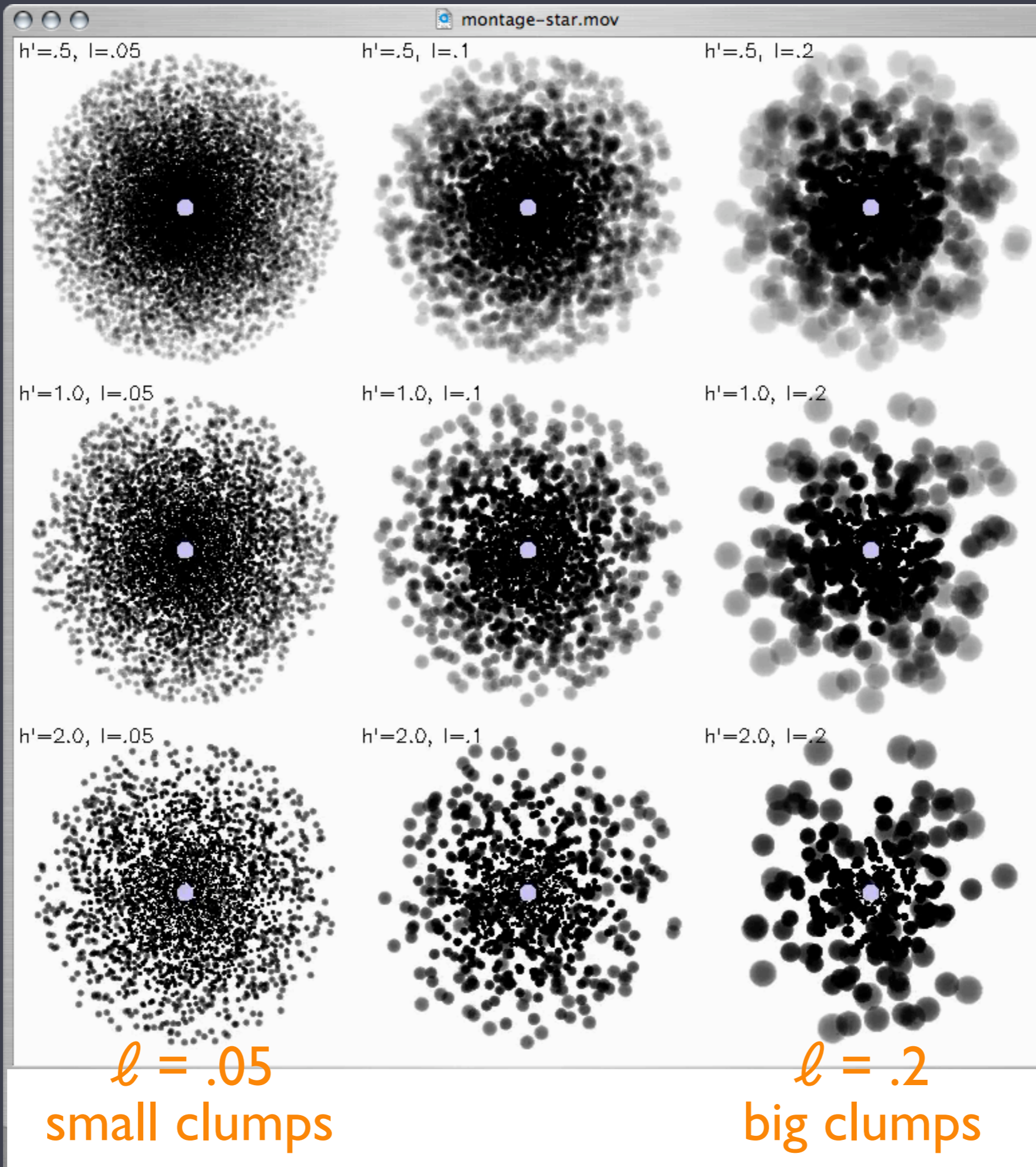
we compare these predictions to data



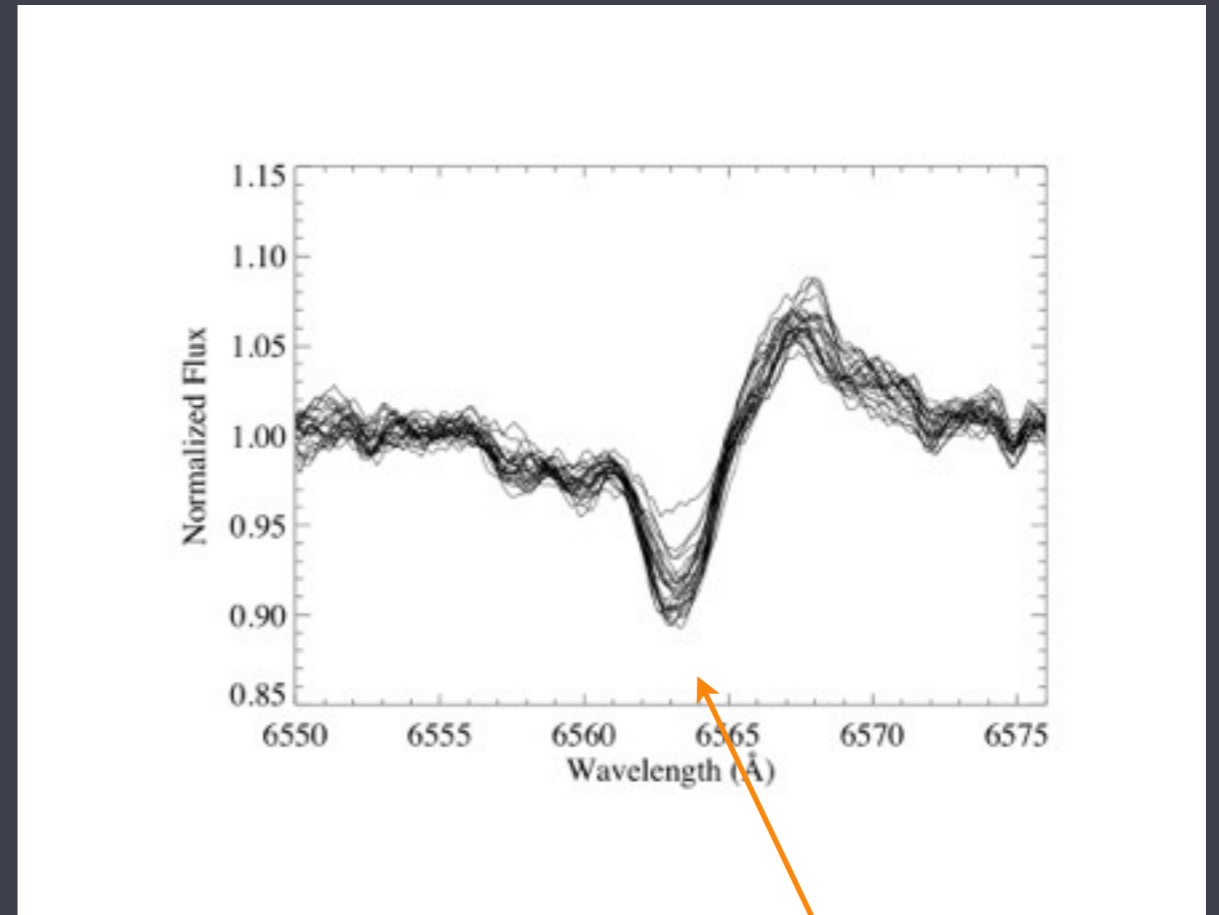
sometimes we contemplate X-ray propagation
through a clumpy and porous medium

less
porous

more
porous



sometimes we make complementary observations
with our telescope on the roof



hydrogen abs/em spectral line
in zeta Ori



Two outstanding questions

1. How do massive stars produce their strong X-ray emission?
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to try to answer these questions

more information on my website

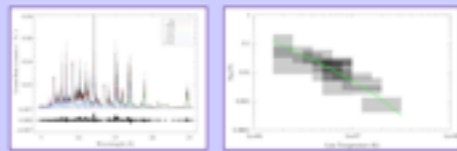
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To get a sense of what the student assistants that assist me do for research

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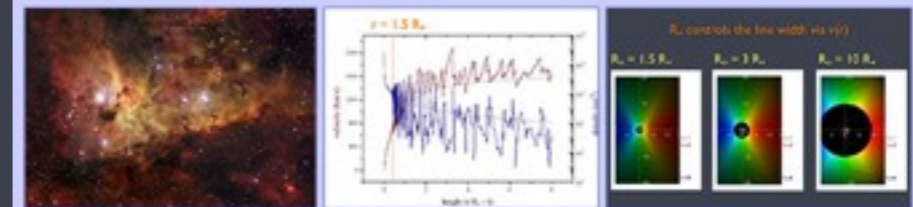
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